NOAO Observing Proposal *Standard proposal* **Panel:** *For office use.*
Date: September 30, 2002 **Category:** Cosmology

High Redshit Supernova Host Galaxy Images

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Abstract of Scientific Justification *(will be made publicly available for accepted proposals)***:**
Type Ia supernovae provide the most direct evidence for an accelerating universe. Continuing studies of supernovae will allow us to improve these measurements, reduce systematic errors, and begin to address the physical nature of the "dark energy" driving the acceleration. To be placed on the Hubble diagram, supernovae require measurements of the peak flux and decay rate of their lightcurves. These measurements in turn require supernova-free host galaxy images. This proposal is for three nights in Spring, 2002 to use WIYN/MIMO to obtain host galaxy images for supernovae discovered in May, 2002, so as to complete observations of these supernovae made during the previous year with Gemini North, the CTIO 4m Blanco telescope, WIYN, and other telescopes.

Summary of observing runs requested for this project

Run	Telescope	Instrument	No. Nights	Moon	Optimal months	Accept. months
1	WIYN	MIMO	3	dark	May - May	May - Jun
2						
3						
4						
5						
6						

Scheduling constraints and non-usable dates *(up to four lines).*

Scientific Justification *Be sure to include overall significance to astronomy. For standard proposals limit text to one page with figures, captions and references on no more than two additional pages.*

The Hubble diagram for Type Ia Supernovae (SNe Ia) provides the most direct current measurement of the expansion history of the universe— and hence the most direct evidence for an accelerating expansion. The Supernova Cosmology Project (SCP) developed an approach to this measurement (Perlmutter *et al.* 1997, 1998, 1999) that resulted in a determination, based on 42 SNe with redshifts between 0.18 and 0.83, of $\Omega_M \sim 0.3$ for a flat universe (Perlmutter *et al.* 1999; see also Riess *et al.* 1998); balloon-based CMB measurements have strengthened the evidence for an accelerating universe by indicating that the geometry of the universe is flat, ruling out the possibility of a low-density open universe. Figure 1 shows current preliminary results, including the Supernovae of Perlmutter *et al.* 1999, and an additional 11 supernovae observed with WFPC-2 on HST between 1998 and 2001, a paper on which is in the final stages of preparation (Knop, 2002).

Further measurements of high-redshift supernovae hold promise for two things. First, over the next few years it will be possible to begin to address the properties of the "dark energy" which is driving the expansion of the universe, by measuring w, the equation of state (p/ρ); a value of $w = -1$ would indicate vacuum energy, or a cosmological constant. In a few years, sufficient supernova data at a range of redshifts together with independent constraints on Ω_M from, e.g. large scale structure, it will be possible to measure w to approximately five times better accuracy than can be done today.

Second, more and better measurements of supernovae will allow us to minimize the effects of systematic errors on the results. For example, the effects of host galaxy dust extinction may be addressed with high-quality color measurements of a large quantity of supernovae. The effects of "grey dust" are addressed by obtaining measurements of supernovae at intermediate ($z < 1$) and high ($z > 1$) redshifts, thereby looking into the epochs of both acceleration and deceleration.

This Proposal: The cosmological measurements described above will build on existing data, and will require the building up of statistics over time. In this proposal, we request the time necessary for SNe discovered during the Spring semester of 2002 with the CTIO 4m telescope to contribute to the cosmological measurements. With the time requested, we will obtain the necessary host galaxy images for supernovae discovered with the CTIO 4m telescope last spring. These images of the host galaxy, uncontaminated by supernova light, will enable an accurate determination of the supernova lightcurve parameters (peak magnitude and decay rate), as described in the "Experimental Design" section. The SNe may then be placed in the Hubble Diagram, and used to contribute to measurements of cosmological parameters and of systematic errors.

References

Knop *et al.*, 2002, BAAS, 199, 1607 (to be submitted to ApJ)
Perlmutter, S., *et al.*, 1997, ApJ, 483, 565
Perlmutter, S., *et al.*, 1998, Nature, 391, 51
Perlmutter, S., *et al.*, 1999, ApJ, 517, 565
Riess, A., *et al.*, 1998, AJ, 116, 1009

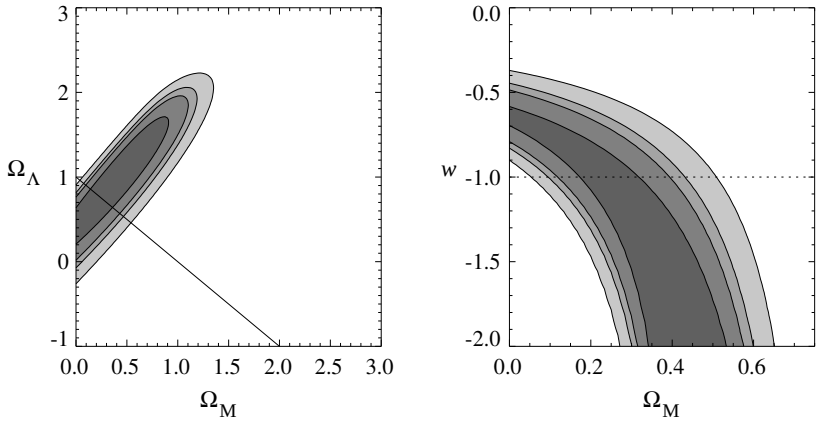

Figure 1: Preliminary limits on cosmological parameters, including data from Perlmutter (1999) plus 11 supernovae observed between 1998 and 2001 with WFPC-2 on HST. On the left are joint limits on Ω_M and Ω_Λ. On the right are joint limits on Ω_M and the equation of state w of the dark energy, under the assumption of a flat universe (motivated by CMB results). In both plots, confidence intervals plotted are 68%, 90%, 95%, and 99%.

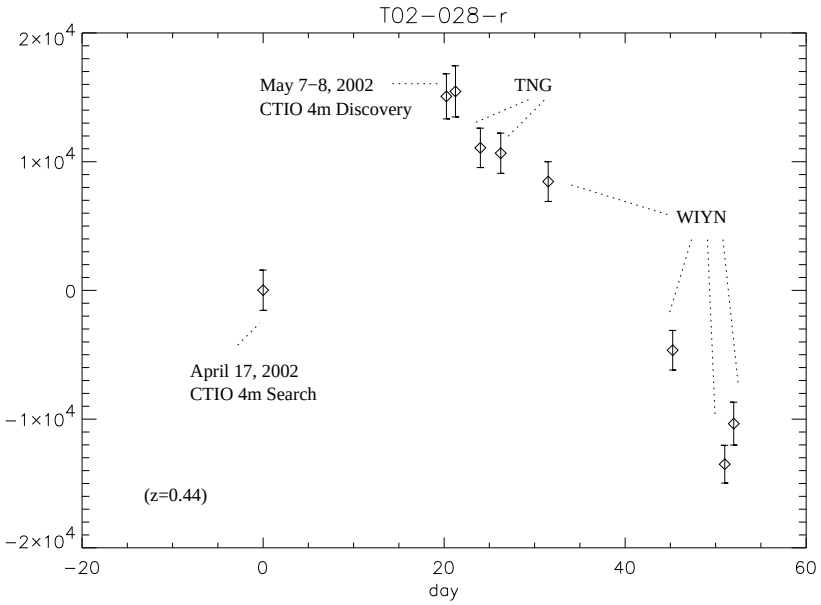

Figure 2: Preliminary R-band lightcurve of supernova T02-028 at z=0.44. Clearly, this supernova had already exploded by the time of the April, 2002 reference run; although the zero level of the flux has been set to the supernova flux on that date, a host galaxy image will allow a true determination of the zero level, plus a better subtraction from the other points (some of whose uncertainties are currently dominated by the uncertainty in the poor reference used).

Experimental Design *Describe your overall observational program. How will these observations contribute toward the accomplishment of the goals outlined in the science justification? If you've requested long-term status, justify why this is necessary for successful completion of the science. (limit text to one page)*

Supernovae were discovered at the CTIO 4m telescope in May 2002 by using a search of two depths. Two adjacent MOSAIC fields (the "deep" fields), plus a third to an intermediate depth, were imaged in the I-band, with the goal of finding supernovae (SNe) at $z > 1$. These fields were observed as long as they were at a favorable airmass. In addition, at the end of the night (after the deep field was too low in the sky), additional "intermediate" fields were imaged in the R-band, with the goal of finding another ~ 6 SNe at redshifts $z < 0.6$. The search technique employed the classical method which the SCP has used successfully over the last several years. The fields were imaged in April, and again ~ 3 weeks later in May. Images from the "search run" in May were convolved to match the point spread function of the images from the "reference run" in April, and subtracted. Residual signals were measured spectroscopically at Keck, Gemini, and VLT to determine supernova identification and to obtain a redshift.

Although poor seeing in April limited our ability to find $z > 1$ candidates, the search was successful and did find SNe. One SN was discovered at $z = 0.912$, and followed with the HST. In addition, five more SNe at $0.2 < z < 0.5$ were discovered (see the table in the Technical Description section). These latter SNe were followed with ground based telescopes through May and June, including with the WIYN telescope during five nights in three different runs. In addition, a SN at $z = 0.35$ discovered simultaneously at Subaru was followed during these ground-based runs.

Although a template lightcurve may be fit to the data in the absence of a host galaxy, this increases both systematic and statistical errors. As the seeing varies between different images of the same SNe, a different amount of host galaxy contaminates the aperture on each image, and hence a single "zero" level in the template fit will not truly be correct for each image. (At the redshifts in question, sizes of host galaxies and separations between host galaxy cores and SNe are comparable in size to photometric apertures of $\sim 1''$.) In addition, any peculiarities or deviations from a SN Ia lightcurve may be hidden by the addition of another parameter to the fit. A host galaxy image fixes the zero level of the lightcurve; it may be convolved to the image quality of each image on the SN lightcurve, allowing an accurate subtraction of the background galaxy light. The images "reference images" from April, 2002 are not sufficient, because they may actually have been taken after the supernova exploded, and have some light from the early rising part of the lightcurve. For the one SN at $z = 0.912$, we have deep images from a year earlier that serve as galaxy references. For the remaining SNe, images taken a year after discovery are necessary. Figure 2 shows a preliminary lightcurve for one of the supernovae discovered with the CTIO 4m telescope in May, 2002.

Because the final reference image is subtracted from each lightcurve point, it must be deeper than the best lightcurve point so that the subtraction does not overmuch degrade the quality of the original data. In addition, the image quality of the data should be no worse than the best seeing from all of the lightcurve images. (The final reference images are still necessarily, if not optimal, if the seeing is worse.) The best image quality on the lightcurves in question come from the WIYN data obtained in May and June of 2002, with seeings of $\sim 0.8''$. Hence, we request time on the WIYN telescope during the Spring of 2003 to obtain the necessary final reference images for these six SNe at $0.2 < z < 0.5$ so that the observations taken with NOAO and other telescopes during Spring 2002 might be used for cosmology.

Use of Other Facilities *Describe how the proposed observations complement data from non-NOAO facilities. For each of these other facilities, indicate the nature of the observations (yours or those of others), and describe the importance of the observations proposed here in the context of the entire program.*

The observations proposed here will naturally complement all of the observations of these supernovae taken during the Spring 2002 semester. In addition to the search at CTIO and the imaging run from WIYN, this includes imaging runs at the WHT and TNG telescopes at the La Palma observatory which obtained additional lightcurve points in May and June. Spectroscopy of these supernovae were obtained with the Gemini North and VLT 8m telescopes. The acquisition of final reference images will allow an accurate determination of the supernova lightcurve parameters, and as such will enable the cosmological application of all of these existing observations.

In addition to the specific application to the lightcurves of the supernovae on the target list, the observations will contribute to a larger project of placing SN Ia on the Hubble diagram to improve measurements of Ω_M and Ω_Λ, as well as the goal of increasing the statistics available for performing systematic studies. The SCP has made observations at a large number of observatories, including those already mentioned, the Keck 10m telescope, and the Hubble Space Telescope, and others in pursuing this survey.

Previous Use of NOAO Facilities *List allocations of telescope time on facilities available through NOAO to the PI during the past 2 years, together with the current status of the data (cite publications where appropriate). Mark with an asterisk those allocations of time related to the current proposal. Please include original proposal semesters and ID numbers when available.*

In the Spring of 2001, the Supernova Cosmology Project (SCP) discovered 17 supernovae (SNe) including two at $z \sim 1.1$ during a search at CTIO (see IAUC 7763 and 7764). These SNe were followed with a variety of ground based telescopes including with MIMO on WIYN, with NIRI on Gemini; the supernovae at the highest redshifts were also observed with HST. Final ground-based reference images were obtained during Spring 2002, and the data analysis is in progress.

In Spring 2002 (NOAO proposal ⋆ 2002A-0330 with Perlmutter as PI), the SCP discovered six SNe with the CTIO 4m telescope, as described in previous sections of this proposal. These supernovae were followed photometrically with WIYN/MIMO, HST, and other telescopes, and spectroscopically with GMOS on Gemini North. The observations proposed here will complete the lightcurves and allow photometric and cosmological analysis of these SNe to proceed.

During the Fall semester of 2002, a joint US/UK proposal (PIs Knop and Hook, program GN-2002B-Q-30) was granted 33.6 hours on the Gemini North telescope, using both GMOS and NIRI. This time will be used to obtain spectroscopic confirmation of supernovae discovered as part of the Subaru Big Project (with discoveries scheduled for the end of October, 2002), as well as for infrared photometry of 1–2 supernovae at redshifts $z > 1$.

Observing Run Details for Run 1: WIYN/MIMO

Technical Description *Describe the observations to be made during this observing run. Justify the specific telescope, the number of nights, the instrument, and the lunar phase. List objects, coordinates, and magnitudes (or surface brightness, if appropriate) in the Target Tables section below (required for WIYN-2hr, WIYN-SYN, YALO, and Gemini runs).*

This run will use the MIMO Camera on the WIYN telescope to obtain final reference images for supernovae lightcurves. Each of these candidates have photometry with host galaxy plus supernova light taken during April, May, and June of 2002, with the CTIO 4m, the TNG, the WHT, and the WIYN telescopes. The final reference images must be deeper than any individual previous observation. In addition, it is advantageous if they have an image quality (seeing) that is no worse than the seeing of the lightcurve images, as described in previous sections. The best seeing images obtained of these supernovae were with MIMO on WIYN, with seeing of $0.7''$–$1''$.

The targets, their coordinates, their redshifts, and exposure time necessary to complete the lightcurve of each SN are listed in the table below. Exposure times are based on existing observations of these targets; previous observations were made during the dark half of the lunar phase, and we are requesting the same phase for the final reference images. Magnitudes listed in the table are *discovery magnitudes* for the supernovae candidates. They are estimates, as the absence of a final reference currently prevents an accurate discovery magnitude measurement. Additionally, many of the previous observation of the SNe were 1–2 magnitudes dimmer, as the lightcurve had faded; it was at these later epochs that the longest exposure times were acquired during Spring 2002. The SNe will have faded completely from view by the time of the run. The final reference images obtained will be psf-matched and subtracted from the lightcurve images, following the procedure used extensively in the past by the Supernova Cosmology Project.

Table: Targets

SN	**RA (2000)**	**Dec (2000)**	z	**Discovery Magnitude**	**Exp. R** (min)	**Exp. I** (min)
S02-064	13:25:46.6	+27:34:20.9	0.35	22.8	90min	120min
T02-028	15:36:25.5	+09:28:18.2	0.44	22.9	90min	120min
T02-047	15:36:29.9	+09:38:42.9	0.49	22.9	90min	120min
T02-029	15:37:07.5	+09:36:18.7	0.21	21.3	60min	90min
T02-030	15:43:24.4	+07:53:57.8	0.51	23.3	90min	120min
T02-037	15:46:57.0	+08:12:36.6	0.49	22.3	90min	120min

Although good image quality helps, there is no benefit to obtaining images with *better* image quality than what was obtained for the supernovae lightcurve points, since the final reference images of the host galaxy will be convolved out to the seeings of those images. As such, MIMO alone on WIYN will be able to obtain images which are ideal for their intended purposes. The relatively large field of view of MIMO is also useful, as it provides a substantial number of field objects which may be used as secondary standard stars for relative photometry (making the data useful even if conditions are not strictly photometric).

If, during the observing run, the seeing is marginal ($> 1''$), we will consider the option of using the WIYN Tip-Tilt Module (WTTM) to obtain images of improved seeing. The NOAO newsletter of September, 2002 suggests that switching between MIMO and the tip-tilt camera is a relatively fast operation, and may be performed repeatedly during an observing run. As such, we will allow the conditions during the run to determine which is the camera that will best serve the goals of the project. Exposure times we quote in the target table are given using the baseline assumption

that MIMO images will be the best choice. In the event that we use WTTM for the images, the slightly lower efficiency (if such still exists during the run) will affect the depth of the final reference images. Although they will then not be as good, they will still be useful for completing the supernovae lightcurves with slightly higher statistical errors.

All told, the observations requested will require three nights on the WIYN telescope. Since most of the targets are at similar Right Ascensions, we will make best use of the telescope with an observing run in May when the targets are at a favorable airmass throughout the night. An observing run during the dark half of the moon's phase in June will also be usable, though not as efficient.

Instrument Configuration

Filters: R, I	Slit:	Fiber cable:
Grating/grism:	Multislit:	Corrector:
Order:	$\lambda_{\rm start}$:	Collimator:
Cross disperser:	$\lambda_{\rm end}$:	Atmos. disp. corr.:

R.A. range of principal targets (hours): 13 to 16

Dec. range of principal targets (degrees): 7 to 27

Special Instrument Requirements *Describe briefly any special or non-standard usage of instrumentation.*

If available, we may switch between MIMO and WTTM during the run.