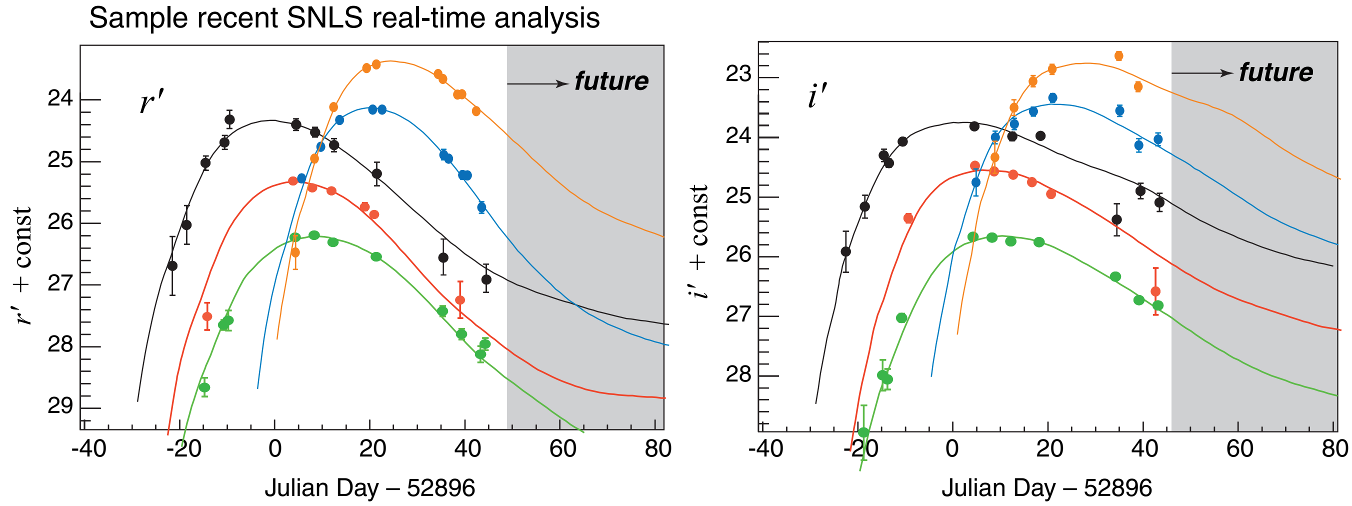
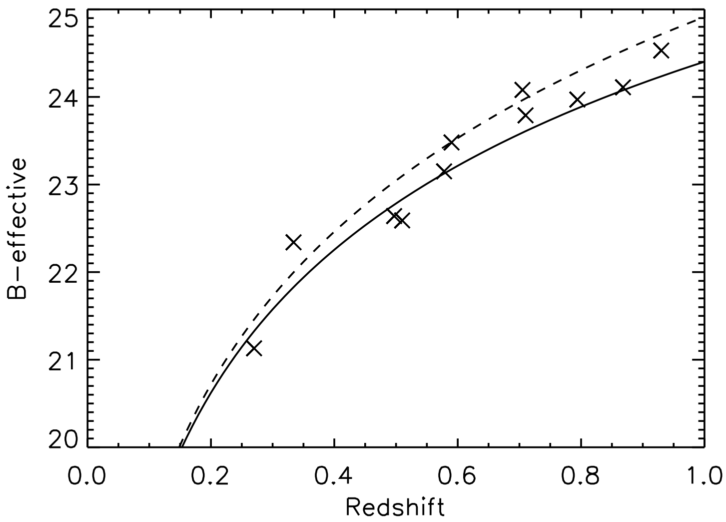


Fig. 1: Example "real-time" light-curves from the survey in r' and i'. At any one epoch, there are always SNe close to maximum light – the most efficient time for spectroscopic follow-up. These candidates are then selected based upon their light-curves, colour and similarity to SN Ia, and sent to Gemini (and other telescopes) for spectroscopic follow-up.

Fig. 2a: A very preliminary Hubble diagram based on SNLS SNe Ia from the first half of semester 2003B and real-time light-curves, showing corrected apparent magntiude versus redshift. The solid and dashed lines show the predicted brightnesses in Einstein-De Sitter and Λ-CDM Universes respectively. The plot demonstrates how Gemini fills the crucial role of populating the high-redshift end of the diagram. Note that this is very preliminary "real-time" data – these SNe have not had either final calibrations applied (they are not yet avalable from CADC) or "final reference" host subtraction (this requires reimaging of the field 1 year later, after the SN has faded).



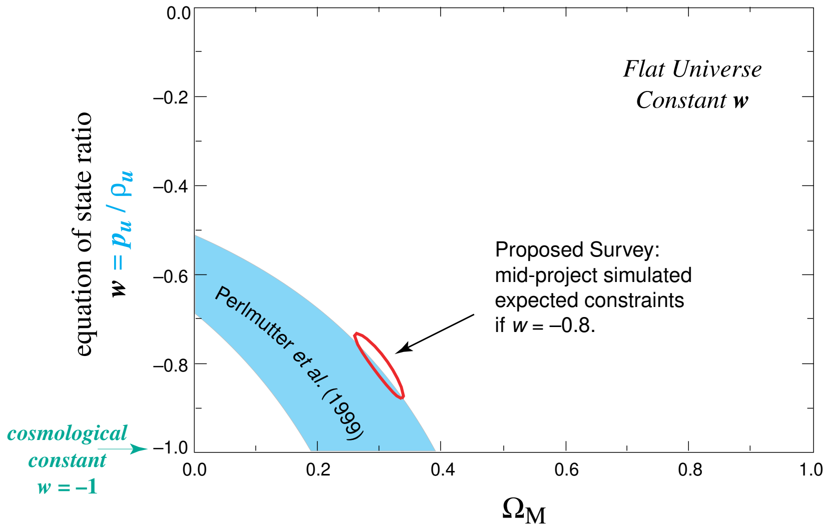


Fig. 2b: Confidence regions in w-Ω_m plane from 300 SNe Ia – corresponding to SNLS mid-survey. This demonstrates the ability of the survey to test whether a pure cosmological constant fits the data, or whether some other model of dark energy is required.

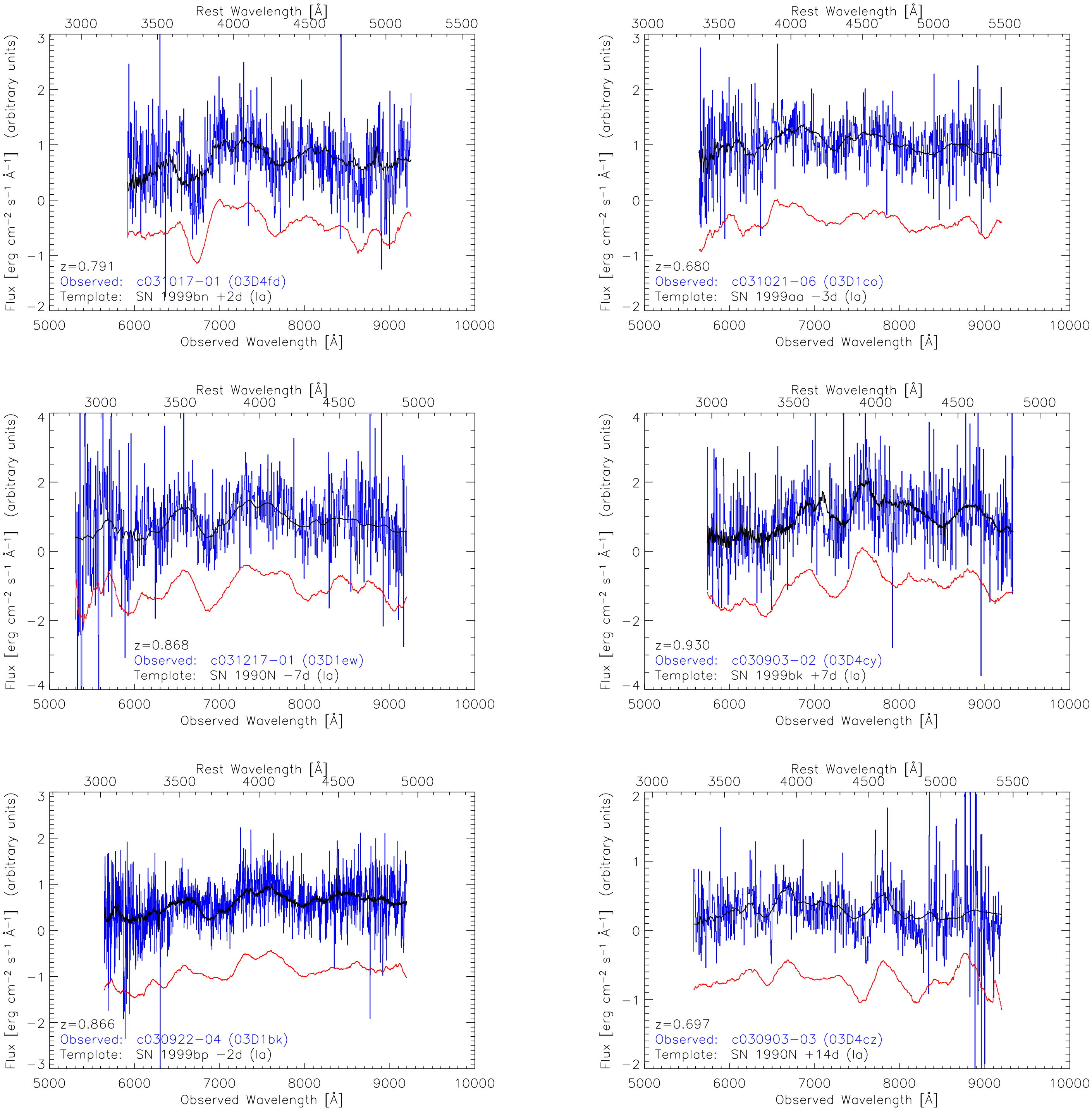


Fig. 3: Example spectra taken during the 2003B GMOS-N time showing the power of nod and shuffle to identify target SNe from $z \simeq 0.7 - 0.9$. Every SN except c030903-03 (bottom right panel) was observed with nod and shuffle. The data is plotted in blue, rebinned to 5Å. Overplotted in black is the best fit template spectrum, in each case a Type Ia. Below (red), we show the data after Savitzky-Golay smoothing. The superior sky subtraction achievable with N+S makes smoothing or rebinning meaningful and enables more secure SN identifications.