*Astronomy & Astrophysics* manuscript no. January 25, 2005
(DOI: will be inserted by hand later)

# Restframe *I*-band Hubble diagram for type Ia supernovae up to $z \sim 0.5^{\star}$

S. Nobili[1,2], R. Amanullah[1], G. Garavini[1,2], A. Goobar[1], C. Lidman[3], V. Stanishev[1], G. Aldering[4], P. Antilogus[2], P. Astier[2], M. S. Burns[5], A. Conley[4,6], S. E. Deustua[7], R. Ellis[8], S. Fabbro[9], V. Fadeyev[4], G. Folatelli[1], R. Gibbons[4], G. Goldhaber[4,6], D. E. Groom[4], I. Hook[10], D. A. Howell[4], A. G. Kim[4], R. A. Knop[11], P. E. Nugent[4], R. Pain[2], S. Perlmutter[4,6], R. Quimby[4], J. Raux[2], N. Regnault[2,4], P. Ruiz-Lapuente[12], G. Sainton[2], K. Schahmaneche[2], E. Smith[11], A. L. Spadafora[4], R. C. Thomas[4], L. Wang[4], and
(THE SUPERNOVA COSMOLOGY PROJECT)

[1] Department of Physics, Stockholm University, Albanova University Center, S-106 91 Stockholm, Sweden
[2] LPNHE, CNRS-IN2P3, University of Paris VI & VII, Paris, France
[3] European Southern Observatory, Alonso de Cordova 3107, Vitacura, Casilla 19001, Santiago 19, Chile
[4] E. O. Lawrence Berkeley National Laboratory, 1 Cyclotron Rd., Berkeley, CA 94720, USA
[5] Colorado College, 14 East Cache La Poudre St., Colorado Springs, CO 80903
[6] Department of Physics, University of California Berkeley, Berkeley, 94720-7300 CA, USA
[7] American Astronomical Society, 2000 Florida Ave, NW, Suite 400, Washington, DC, 20009 USA
[8] California Institute of Technology, E. California Blvd, Pasadena, CA 91125, USA
[9] CENTRA-Centro M. de Astrofísica and Department of Physics, IST, Lisbon, Portugal
[10] Department of Physics, University of Oxford, Nuclear & Astrophysics Laboratory, Keble Road, Oxford, OX1 3RH, UK
[11] Department of Physics and Astronomy, Vanderbilt University, Nashville, TN 37240, USA
[12] Department of Astronomy, University of Barcelona, Barcelona, Spain

Received ...; accepted ...

**Abstract.** Using a novel technique for fitting restframe *I*-band light curves of 42 type Ia supernovae, we construct a Hubble diagram including 26 SNe with $0.01 < z < 0.1$. Adding two SNe at $z \sim 0.5$ yields results consistent with a flat $\Lambda$-dominated "concordance universe" $(\Omega_M, \Omega_\Lambda)$=(0.25,0.75). For one of these, SN 2000fr, new near infrared data are presented. The high redshift supernova NIR data are also used to test for systematic effects in the use of SNIa as distance estimators. A flat, $\Lambda = 0$, universe where the faintness of supernovae at $z \sim 0.5$ is due to grey dust homogeneously distributed in the intergalactic medium is disfavoured based on the high-z Hubble diagram using this small data-set. However, the uncertainties are large and no firm conclusion may be drawn. We explore the possibility of setting limits on intergalactic dust based on $B - I$ and $B - V$ colour measurements, and conclude that about 20 well measured SNe are needed to have statistical significance of the results. We also show that the high redshift restframe *I*-band data points are better fit by light curve templates that show a prominent second peak, suggesting that they are not intrinsically underluminous.

**Key words.** cosmology:observations, supernovae: general

## 1. Introduction

Observations of type Ia supernovae (SNe) in restframe *B*-band at redshifts of $z \sim 0.5$ and above have shown that they are best fit by a cosmological model that includes a cosmological constant or some other form of dark energy (Perlmutter et al., 1998; Garnavich et al., 1998; Riess et al., 1998; Schmidt et al., 1998; Perlmutter et al., 1999; Tonry et al., 2003; Knop et al., 2003; Riess et al., 2004). The evidence for dark energy is supported by cross-cutting cosmological results, such as the measurement of the cosmic microwave background anisotropy, which indicates a flat universe (De Bernardis et al., 2000; Jaffe et al., 2001; Sievers et al., 2003; Spergel et al., 2003); the evo-

*Send offprint requests to*: S. Nobili , serena@physto.se

⋆ Based in part on observations taken at the European Southern Observatory using the ESO Very Large Telescope on Cerro Paranal (ESO Program 265.A-5721(B)). Based in part on observations made with the NASA/ESA Hubble Space Telescope, obtained at the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5-26555. These observations are associated with program GO-8346. Based in part on data collected from the Canada-France-Hawaii Telescope Corporation, which is operated by the National Research Council of Canada, le Centre National de la Recherche Scientifique de France, and the University of Hawaii.

lution in the number density of X-ray emitting galaxy clusters (Borgani et al., 2001; Henry, 2001; Schuecker et al., 2003) and galaxy redshift surveys (Efstathiou et al., 2002), which indicate that $\Omega_M \approx 0.3$. Taken together, these independent measurements suggest a concordance universe ($\Omega_M$,$\Omega_\Lambda$)≅(0.25,0.75). However, the SN Hubble diagram remains the most direct approach currently in use for studying cosmic acceleration, and, thus, possible systematic effects affecting the observed brightness of type Ia supernovae should be carefully considered. Some of these have already been addressed in previous SCP publications (Perlmutter et al., 1997, 1998; Sullivan et al., 2003; Knop et al., 2003), such as uncorrected host galaxy extinction (see e.g. Rowan-Robinson (2002)) and dimming by photon-axion mixing over cosmological distances (Csaki et al., 2002; Deffayet et al., 2001; Mörtsell et al., 2002; Östman & Mörtsell, 2004), dimming by intergalactic grey dust (Aguirre, 1999a,b; Mörtsell & Goobar, 2003).

Determining cosmological distances through type Ia supernova fluxes at longer restframe wavelengths offers potential advantages, e.g. less sensitivity to dust along the line of sight, either in the host galaxy or in the intergalactic medium. On the other hand, the "standard candle" properties at these wavelengths and the possibility of additional systematic effects need to be investigated.

Using restframe *I*-band, the uncertainty due to extinction by dust is greatly reduced as compared with restframe *B*-band measurements. For example, for Milky-Way type dust ($R_V \sim 3$) the ratio of extinction for the two bands is sizable, $A_B/A_I \sim 2-3$. In general, the extinction corrections become less dependent on our knowledge of the intrinsic supernova colours and dust properties.

*I*-band light curves typically show a second peak 15-30 days after the first maximum. It has been suggested that the intensity and time-difference between the first and second *I*-band peaks are related to the intrinsic luminosity of the type Ia SNe, appearing later and more evident for normal type Ia and earlier and fainter for underluminous ones (Hamuy et al., 1996a; Wang et al., 2005). Thus, building *I*-band light curves for type Ia supernovae offers the possibility of probing brightness evolution.

The scope of this work is to test the feasibility of using restframe *I*-band for cosmological distance measurements, using data available to date, and assess the importance of observing in this wavelength range for future samples of SNe. For that purpose, we develop a template fitting technique to estimate the first ($I_{max}$) and second ($I_{sec}$) *I* light curve peaks, which we apply to 42 nearby SNe Ia. $I_{max}$ is then used to build a SN Ia Hubble diagram reaching out to $z \sim 0.5$. New Infrared data of a supernova at redshift $z = 0.543$, SN 2000fr, are presented and used in the Hubble diagram, together with two other SNe in the same redshift range, available in literature: SN 1999Q (Riess et al., 2000) and SN 1999ff (Tonry et al., 2003). The properties of the second peak in the restframe *I*-band light curves are investigated. Furthermore, additional colour information is used to test for extinction by non-conventional dust for the $z \sim 0.5$ supernovae, **similar to Riess et al. (2000).**

## 2. I-band light curve fitting

The second light curve peak seen in nearby type Ia SNe varies in strength and position with respect to the primary maximum and complicates the use of a single parametrised template for light curve fitting, which is currently often applied for *B* and *V*-band, (see e.g. Perlmutter et al. (1997); Goldhaber et al. (2001) for an example of the timescale stretch factor approach).

**We have developed a method for fitting *I*-band light curves using five free parameters to modify one template[1]. Our fitting procedure can be summarised as follows: one template is used to fit the time ($t_1$) and the first peak magnitude ($I_1$), together with a stretch factor ($s_I$), which is also applied to the second template shifted in time to fit the time ($t_2$) and magnitude of the second peak ($I_2$). The underlying function is**

$$I = I_1\mathcal{T}(s_I(t-t_1)) + I_2\mathcal{T}(s_I(t-t_2))$$

**where $\mathcal{T}$ is the template used. The five parameters fitted are thus: $\{t_1, t_2, I_1, I_2, s_I\}$, (see Table 1).** A similar approach is also proposed by Wang et al. (2005) who calls it "super-stretch" to emphasise its extension of the stretch approach.

Contardo et al. (2000) proposed a model composed of as many as 4 functions for a total of 10 parameters in order to fit all *UBVRI*-bands. Their method used two Gaussian functions to fit the two peaks together with a straight line to fit the late time decline and an exponential factor for the pre-max rising part of the light curve. **In this way is possible to describe type Ia SNe light curves over a wide range of epochs and in all optical bands, though, as the authors recognise, does not represent accurately the second peak in the *I*-band due to the influence of the linear decline. However, the main disadvantage of their method for our purpose, is the large number of free parameters, which requires that the light curve for the object to be extremely well sampled.**

The use of the template in place of the Gaussian function reduces the number of free parameters by a factor of two. Moreover, with this choice, no additional functions are needed to fit the pre-max rising part of the light curve nor the late time decline. **Implicitly, we have thus assumed that the rising part of the light curve in *I*-band is the same as in the template used. As we will see, this assumption is not always optimal.** As our goal is only to measure the position and amplitude of the first two peaks, we limit the fit to 40 days after maximum, neglecting the late time decline. Note that, unless otherwise specified, the supernova phase always refers to the time relative to restframe *B*-band light curve maximum.

### 2.1. The data set

We applied this method to fit a sample of local SNe for which both *B* and *I*-band data are available in the literature. The SNe primarily comes from the Calan/Tololo (Hamuy et al., 1996a), CfA (Riess et al., 1999) and CfA2 (Jha, 2002) data sets. Data

[1] **The *B*-band template in Nugent et al. (2002) has been used, although there are not physical reasons for choosing the *B*-band over other bands or other kind of templates.**

| | |
|---|---|
| $t_1$ | time of the peak of the first $\mathcal{B}$ template |
| $I_1$ | peak magnitude of the first $\mathcal{B}$ template |
| $t_2$ | time of the peak of the second $\mathcal{B}$ template |
| $I_2$ | peak magnitude of the second $\mathcal{B}$ template |
| $s_I$ | stretch factor of the time axis |
| $t_{\max}$ | time of the first $I$ light curve peak |
| $I_{\max}$ | first $I$ light curve peak magnitude |
| $t_{\sec}$ | time of the second $I$ light curve peak |
| $I_{\sec}$ | second $I$ light curve peak magnitude |

Table 1 Summary of the parameters used in this work to describe the $I$-band light curve. The first five parameters are determined by fitting the data (see text for details). The next four parameters are determined from the first set and are the actual time and peak values of the light curve.

from three other well studied individual supernovae were also included: SN 1989B (Wells et al., 1994), SN 1994D (Richmond et al., 1995) and one underluminous supernova: SN 1991bg. There are at least two available data sets in restframe $I$-band for SN 1991bg, one published by Filippenko et al. (1992b) with quite good coverage from about 3 days after $B$-band maximum light to +60 days, another one published by Leibundgut et al. (1993) with four data points, the first of which is at the time of $B$-band maximum. The agreement between the two data sets was assessed by comparing the measurements taken at the same date, i.e. JD=2448607, where we found a difference of 0.06 mag. We take this as an estimate of the measurement uncertainty in data of Leibundgut et al. (1993) as no uncertainties are reported in that work.

### 2.2. Fitting method and results

Only supernovae with at least 6 $I$-band data points and covering in time constraining both peaks were selected for light curve fitting. This resulted in a total of 42 SNe. Table 2 lists the parameters resulting from the fitting procedure. **Note that the fit was performed in units of flux, thus, the error bars were considered symmetric in flux and treated as gaussian. The parameters given in the table are transformed to magnitudes.**

Prior to fitting, all data points were $K$-corrected to restframe $I$-band as in (Kim et al., 1996; Nugent et al., 2002), assuming a Bessell $I$-band filter transmission curve (Bessell, 1990) and time information from the available $B$-band data. A new spectroscopic template, which is a slightly modified version of the template found in Nugent et al. (2002), was built for computing the $K$-corrections. We have preserved the SED from the UV through the Si II 6150Å feature, following Nobili et al. (2003), but red ward of that we have incorporated additional spectra from the Supernova Cosmology Project (SCP) Spring 1999 search (Aldering, 2000; Nugent & Aldering, 2000) to improve this region as the original template was sparse and required a lot of interpolation. A potential source of systematic uncertainties in the $K$-corrections is due to the wide Ca IR triplet absorption feature, found to vary considerably among Type Ia supernovae (Strolger et al., 2002; Nugent et al., 2002). We have estimated this systematic uncertainty as a function of redshift for $0.01 < z < 0.1$ by computing the $K$-correction for diverse nearby SNe Ia at different epochs. They show a dispersion of up to 0.05 mag for increasing redshift. We take this as a conservative estimate of the uncertainty in all $K$-corrections in this wavelength region.

Note that the values of $I_2$ reported in Table 2 are not the actual magnitudes of the secondary peak , $I_{\sec}$, but a parameter indicating the size of the contribution of the secondary B-band template to the overall I-band template.

In figure 1, all the fitted light curves are shown. They are sorted in chronological order, except for the two very underluminous supernovae: SN 1991bg and SN 1997cn, displayed at the bottom of the figure. As the date of $B$-band maximum for SN 1997cn is unknown, the origin of the time axis was set to the epoch ($JD = 2450597.75$) when this supernova was first observed. Note that the second peak for the underluminous supernovae is almost completely absent, resulting in a value of $I_2 \sim 2.5$ to 3 magnitudes fainter than $I_1$.

Our sample includes SNe that are classified as spectroscopically peculiar, showing similarities with the over-luminous SN 1991T **(Filippenko et al., 1992a; Ruiz-Lapuente et al., 1992; Phillips et al., 1992)**. These are SN 1995bd, SN 1997br, SN 1998ab, SN 1998es, SN 1999aa, SN 1999ac, SN 1999dq and SN 1999gp**(Li et al., 2001; Howell, 2001; Garavini et al., 2004b)**. One supernova, SN 1993H, was reported to show similarities with the spectrum of the peculiar underluminous 1986G (Hamuy et al., 1993). **However, as we will see in this work, we do not find all of these to show peculiarities in their *I*-band light curve shape when compared to spectroscopically normal SNe Ia.** Recently, Krisciunas et al. (2003), built the Hubble diagram for SNe Ia in infrared $J, H$ and $K$-bands out to $z = 0.04$, and reported that three spectroscopically peculiar SNe, SN 1999aa, SN 1999ac and SN 1999aw, do not show a behaviour different than that of normal SNe. **In our attempt to assess a greater homogeneity of SNe as standard candles in the *I*-band than in *B*-band, we choose not to exclude peculiar SNe from our sample, and instead monitor possible deviant behaviour of these objects.**

Analysing the results of our fits, we found that type Ia SNe show a variety of properties for the $I$-band light curve shape. In particular we noticed that the light curve could peak between −3 days and +3 days w.r.t. $B_{\max}$, as shown in figure 2 (left panel). The time of the second peak, $t_{\sec}$ (relative to $B_{\max}$), is shown in the right panel. The distribution of $t_{\max}$ is centred at day -0.3 and has a dispersion of $\sigma = 1.3$ days. $t_{\sec}$ is centred at 23.6 with a dispersion of $\sigma = 4.4$ days. **The result shown in figure 2 is statistically compatible with the one of Contardo et al. (2000), which is based on a smaller sample, although they lead to different conclusions, e.g. Contardo et al. (2000) found that the time of *I*-band maximum occurs earlier than the time of *B*-band maximum**.

The fits have reduced $\chi^2$ values (see Table 2) that are generally around unity, except for a few cases. In particular we note that SN 1994D has a $\chi^2/dof \sim 26$. Although it has been suggested that the uncertainties for this supernova may be underestimated (see Knop et al. (2003)), the trend in the residuals shows that this SN is not well described by the model, indicating the limitations of the fitting template. As in other cases we

find a systematic trend, especially in the rising part of the light curves. Only 13 SNe have pre-maximum data, and 6 out of 13 are not well fitted in their rise. **This is likely to affect the peak magnitude fitted for these 6 SNe more significantly than the fitted time of maximum.** We investigated possible systematic effects in the fitted light curve maximum due to this, but found no evidence of a trend in the residuals in the Hubble diagram (see Section3) for these 6 SNe.

While the $\chi^2$ gives a measurement of the goodness of the fit, in the next section we test the robustness and accuracy of the parameter estimation in our fitting method, reported in Table 2.

### 2.3. Monte-Carlo tests of the fitting method

Given the heterogeneous origin of the data sample, the quality and the sampling of the individual SN light curves vary considerably. Only a few supernovae have excellent time coverage in the *I*-band, resulting in a wide range of accuracy in the fitted parameters. The robustness of the fitting procedure was tested for all circumstances of data quality and time sampling in our sample by means of Monte Carlo simulations. We generated 1000 sets of simulated light curves for each supernova. The synthetic data points had the same time sampling as the real light curves and with deviations from the best fit template randomly drawn from a Gaussian whose width was set by the published uncertainties. The simulated light curves were fitted using the same method as for the experimental data sets. The distribution of the fitted parameters from the simulated data was compared with the input data from the fits of the experimental data. The mean value in the distribution of each parameter generally coincides with that expected, i.e. within one standard deviation. There is no evidence for biased fit parameters. This lends confidence that the fitting procedure is robust, and given the model of the light curve template, will not yield biased estimation of the parameters. However, out of 42 supernovae, 2 cases showed significant deviations; both of these are identified as spectroscopically peculiar SNe,

- SN 1997br: in about 3% of the cases we found $\chi^2/dof \geq 7$. These cause the distributions to have standard deviation larger than the uncertainty computed by the fitting procedure. Excluding these deviant simulated fits results in the expected distributions.
- SN 1998ab: the simulated data sets result in two populations of parameters, which reduces to one when imposing a cut on $\chi^2/dof \leq 6$. Note that about 78% of the simulations satisfy this condition.

In both cases a cut of the tail of the $\chi^2$ distribution was enough to discriminate between the results, thus these supernovae were kept in the remaining analysis since the $\chi^2/dof$ of the fits to the real data fulfilled the cut criteria. **We note, incidentally, that these SNe happen to be the two with the smallest ratios between peak and dip in their light curves. However, the first peak is determined by only one and two data points each. A close look at the simulated light curves indicates the latter to be the most likely cause of the seldom failure of the MC simulation.**

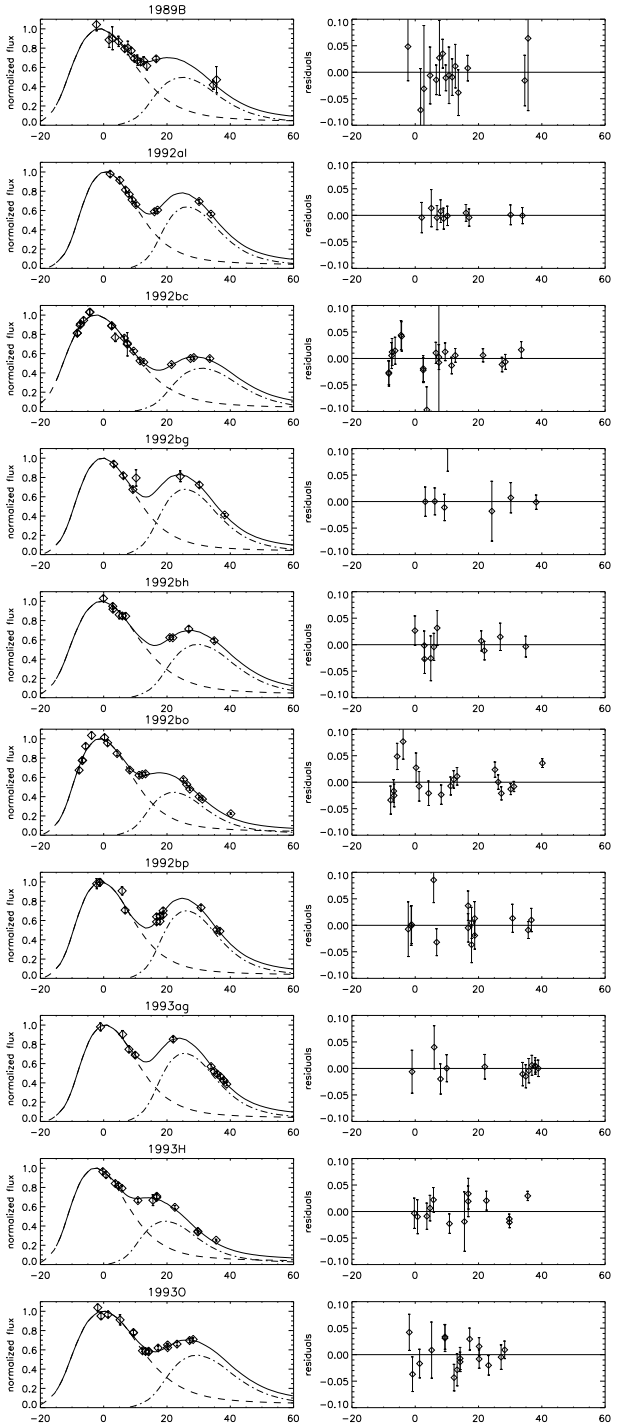


Fig. 1 *I*-band light curve fits. On the ordinate is the flux normalised to the first peak, on the abscissa the restframe time since *B*-band maximum. The dashed line and the dash-dotted line represent the two templates used to fit the first and second peak respectively.

For the rest of the supernovae, the simulations confirmed the expected parameters, giving general confidence in the robustness of the procedure and the accuracy of the uncertainties on the parameters given in Table 2.

| SN | z | $s_I$ | $t_1$ | $I_1$ | $t_2$ | $I_2$ | $\chi^2$ | $dof$ |
|---|---|---|---|---|---|---|---|---|
| 1989B[1] | 0.002 | 1.100 ± 0.126 | -0.590 ± 1.549 | 11.752 ± 0.059 | 23.123 ± 2.079 | 12.513 ± 0.175 | 5.12 | 10 |
| 1991bg[4] | 0.005 | 1.104 ± 0.034 | 3.274 ± 0.380 | 13.521 ± 0.006 | 28.800 ± 1.123 | 16.506 ± 0.122 | 21.78 | 15 |
| 1992al[4] | 0.015 | 0.952 ± 0.054 | 1.176 ± 1.016 | 15.039 ± 0.043 | 27.887 ± 1.078 | 15.527 ± 0.069 | 0.59 | 5 |
| 1992bc[4] | 0.020 | 1.121 ± 0.030 | -1.579 ± 0.146 | 15.639 ± 0.014 | 27.677 ± 0.492 | 16.510 ± 0.026 | 18.78 | 15 |
| 1992bg[4] | 0.035 | 0.963 ± 0.065 | 0.257 ± 1.929 | 17.543 ± 0.086 | 26.392 ± 2.543 | 17.963 ± 0.060 | 3.18 | 2 |
| 1992bh[4] | 0.045 | 1.086 ± 0.145 | -0.008 ± 1.029 | 17.899 ± 0.028 | 26.766 ± 2.494 | 18.543 ± 0.116 | 4.16 | 5 |
| 1992bo[4] | 0.019 | 0.952 ± 0.019 | -0.609 ± 0.162 | 16.064 ± 0.016 | 23.403 ± 0.374 | 16.944 ± 0.042 | 43.91 | 13 |
| 1992bp[4] | 0.079 | 0.891 ± 0.053 | -0.292 ± 0.710 | 18.962 ± 0.029 | 27.750 ± 1.287 | 19.346 ± 0.084 | 10.14 | 9 |
| 1993ag[4] | 0.049 | 0.924 ± 0.058 | 1.092 ± 0.863 | 18.384 ± 0.039 | 26.904 ± 2.010 | 18.759 ± 0.060 | 2.57 | 7 |
| 1993H[4] | 0.024 | 0.953 ± 0.033 | -2.072 ± 0.960 | 16.664 ± 0.040 | 20.556 ± 0.943 | 17.532 ± 0.048 | 23.61 | 8 |
| 1993O[4] | 0.051 | 1.089 ± 0.080 | 0.462 ± 0.917 | 18.197 ± 0.023 | 26.248 ± 1.108 | 18.858 ± 0.034 | 15.97 | 11 |
| 1994ae[3] | 0.004 | 1.051 ± 0.017 | -1.110 ± 0.142 | 13.383 ± 0.018 | 26.664 ± 0.312 | 14.032 ± 0.042 | 22.65 | 15 |
| 1994D[2] | 0.002 | 0.891 ± 0.004 | -1.007 ± 0.043 | 12.177 ± 0.004 | 25.219 ± 0.091 | 12.836 ± 0.008 | 546.17 | 21 |
| 1994M[3] | 0.023 | 0.945 ± 0.041 | 0.035 ± 1.262 | 16.513 ± 0.050 | 24.824 ± 1.027 | 17.139 ± 0.063 | 26.73 | 8 |
| 1994T[3] | 0.035 | 0.740 ± 0.026 | 2.371 ± 1.319 | 17.458 ± 0.049 | 30.187 ± 0.988 | 17.840 ± 0.053 | 9.22 | 3 |
| 1995al[3] | 0.005 | 1.158 ± 0.046 | -1.002 ± 0.472 | 13.526 ± 0.022 | 24.696 ± 0.751 | 14.149 ± 0.057 | 8.72 | 11 |
| 1995bd[3] | 0.016 | 1.166 ± 0.025 | -0.095 ± 0.112 | 16.082 ± 0.012 | 26.516 ± 0.367 | 16.632 ± 0.060 | 20.70 | 11 |
| 1995D[3] | 0.007 | 1.267 ± 0.054 | -1.408 ± 0.682 | 13.708 ± 0.026 | 24.763 ± 0.744 | 14.454 ± 0.037 | 10.34 | 20 |
| 1995E[3] | 0.012 | 1.026 ± 0.040 | 0.067 ± 0.635 | 15.393 ± 0.024 | 26.340 ± 0.950 | 16.093 ± 0.050 | 6.78 | 9 |
| 1996ai[3] | 0.003 | 1.115 ± 0.024 | -2.042 ± 0.495 | 13.986 ± 0.022 | 24.947 ± 0.504 | 14.675 ± 0.023 | 48.38 | 5 |
| 1996bl[3] | 0.036 | 0.942 ± 0.026 | 1.932 ± 0.365 | 17.079 ± 0.022 | 29.619 ± 0.684 | 17.702 ± 0.033 | 9.44 | 4 |
| 1996bo[3] | 0.017 | 1.072 ± 0.010 | -0.325 ± 0.113 | 15.701 ± 0.005 | 24.675 ± 0.186 | 16.253 ± 0.013 | 85.76 | 7 |
| 1996C[3] | 0.030 | 1.059 ± 0.038 | 0.278 ± 1.085 | 16.829 ± 0.052 | 27.206 ± 0.928 | 17.521 ± 0.030 | 23.02 | 6 |
| 1996X[3] | 0.007 | 1.079 ± 0.041 | -1.944 ± 0.371 | 13.399 ± 0.012 | 24.524 ± 0.774 | 14.170 ± 0.033 | 7.87 | 10 |
| 1997bp[5] | 0.008 | 1.235 ± 0.049 | 1.092 ± 0.277 | 14.134 ± 0.006 | 25.847 ± 0.484 | 14.637 ± 0.018 | 8.18 | 6 |
| 1997bq[5] | 0.009 | 1.014 ± 0.014 | 0.479 ± 0.099 | 14.580 ± 0.017 | 25.130 ± 0.230 | 15.139 ± 0.017 | 16.26 | 6 |
| 1997br[5] | 0.007 | 1.349 ± 0.035 | 0.769 ± 0.110 | 13.683 ± 0.020 | 20.672 ± 0.451 | 14.295 ± 0.037 | 13.75 | 5 |
| 1997cn[5] | 0.017 | 0.840 ± 0.049 | -1.711 ± 0.776 | 16.426 ± 0.027 | 24.070 ± 1.067 | 18.255 ± 0.156 | 11.78 | 7 |
| 1997dg[5] | 0.031 | 0.965 ± 0.060 | -1.324 ± 1.082 | 17.240 ± 0.035 | 26.675 ± 2.351 | 17.744 ± 0.051 | 1.18 | 1 |
| 1997e[5] | 0.013 | 0.931 ± 0.031 | -1.375 ± 0.315 | 15.477 ± 0.007 | 23.879 ± 0.386 | 16.064 ± 0.022 | 8.60 | 3 |
| 1998ab[5] | 0.027 | 1.413 ± 0.046 | 0.279 ± 0.179 | 16.485 ± 0.021 | 19.787 ± 0.556 | 17.049 ± 0.042 | 8.40 | 5 |
| 1998dh[5] | 0.009 | 0.997 ± 0.011 | -0.337 ± 0.169 | 14.099 ± 0.015 | 26.025 ± 0.275 | 14.682 ± 0.024 | 0.77 | 1 |
| 1998es[5] | 0.011 | 0.980 ± 0.140 | -2.309 ± 0.511 | 14.083 ± 0.016 | 24.278 ± 1.062 | 14.928 ± 0.085 | 6.90 | 6 |
| 1998v[5] | 0.018 | 0.940 ± 0.036 | 0.545 ± 0.611 | 15.790 ± 0.017 | 25.688 ± 0.698 | 16.100 ± 0.056 | 9.65 | 2 |
| 1999aa[5] | 0.014 | 1.288 ± 0.014 | 0.301 ± 0.072 | 15.242 ± 0.007 | 25.585 ± 0.180 | 15.871 ± 0.027 | 73.43 | 9 |
| 1999ac[5] | 0.009 | 1.210 ± 0.027 | 1.328 ± 0.283 | 14.321 ± 0.006 | 23.045 ± 0.519 | 15.057 ± 0.032 | 11.26 | 7 |
| 1999cl[5] | 0.008 | 0.970 ± 0.118 | -0.107 ± 0.571 | 13.139 ± 0.022 | 22.206 ± 1.647 | 13.688 ± 0.101 | 0.63 | 3 |
| 1999dq[5] | 0.014 | 1.179 ± 0.024 | -0.324 ± 0.126 | 14.785 ± 0.006 | 25.492 ± 0.198 | 15.312 ± 0.013 | 34.19 | 15 |
| 1999gp[5] | 0.027 | 1.261 ± 0.044 | -1.600 ± 0.263 | 16.405 ± 0.008 | 26.011 ± 0.526 | 17.045 ± 0.024 | 7.51 | 6 |
| 2000cn[5] | 0.023 | 0.799 ± 0.026 | -0.362 ± 0.206 | 16.679 ± 0.015 | 23.965 ± 0.441 | 17.256 ± 0.062 | 15.40 | 7 |
| 2000dk[5] | 0.017 | 0.809 ± 0.016 | -1.342 ± 0.177 | 15.767 ± 0.007 | 24.297 ± 0.429 | 16.222 ± 0.048 | 20.33 | 4 |
| 2000fa[5] | 0.021 | 1.119 ± 0.050 | -0.235 ± 0.278 | 16.289 ± 0.037 | 24.911 ± 0.423 | 16.888 ± 0.111 | 0.20 | 2 |

Table 2 Results of the $I$-band light curve fit of 42 nearby supernovae: $t_1$ and $I_1$ are the parameters for the time and amplitude fitted on the first $B$-band template, $t_2$ and $I_2$ are the parameters for the time and amplitude fitted on the second $B$-band template, and $s_I$ is the stretch factor. The data were taken from: [1] Wells et al. (1994); [2] Richmond et al. (1995); [3] Riess et al. (1999); [4] Hamuy et al. (1996a); [5] Jha (2002); [6] Filippenko et al. (1992b); Leibundgut et al. (1993).

### 2.4. Intrinsic variations

We investigated possible relations between $I$-band and $B$-band parameters. Published $B$-band data, when available, were fitted using a template in order to determine the time of maximum luminosity, the stretch factor, $s_B$, and the amplitude of maximum, $m_B$, following Goldhaber et al. (2001). A width-luminosity relation was found for the first $I$-band light curve peak. Figure 3 shows the $I$-band absolute magnitude versus the stretch factor in $B$-band for SNe with $z_{CMB} \geq 0.01$, where the distance (in Mpc) to each SN was calculated from its redshift, assuming a value for the Hubble constant, $H_0 = 72$ km s$^{-1}$ Mpc$^{-1}$. The error bars in figure 3 include an uncertainty of 300 km s$^{-1}$ on the redshifts to account for peculiar velocities of the host galaxies. Corrections for Milky Way and host galaxy extinction were also applied, i.e.

$$M^I_{\max} - 5\log(H_0/72) = I_{\max} - A_I^{MW} - A_I^{host} - 25 - 5\log(d_L)$$

The host galaxy extinction corrections that is applied to most of the supernovae is the weighted average of the three estimates given in Table 2 of Phillips et al. (1999) assuming $R_I = 1.82$. The extinction for the supernovae in the CfA2 data set was calculated following the same procedure, using $B$ and $V$-band

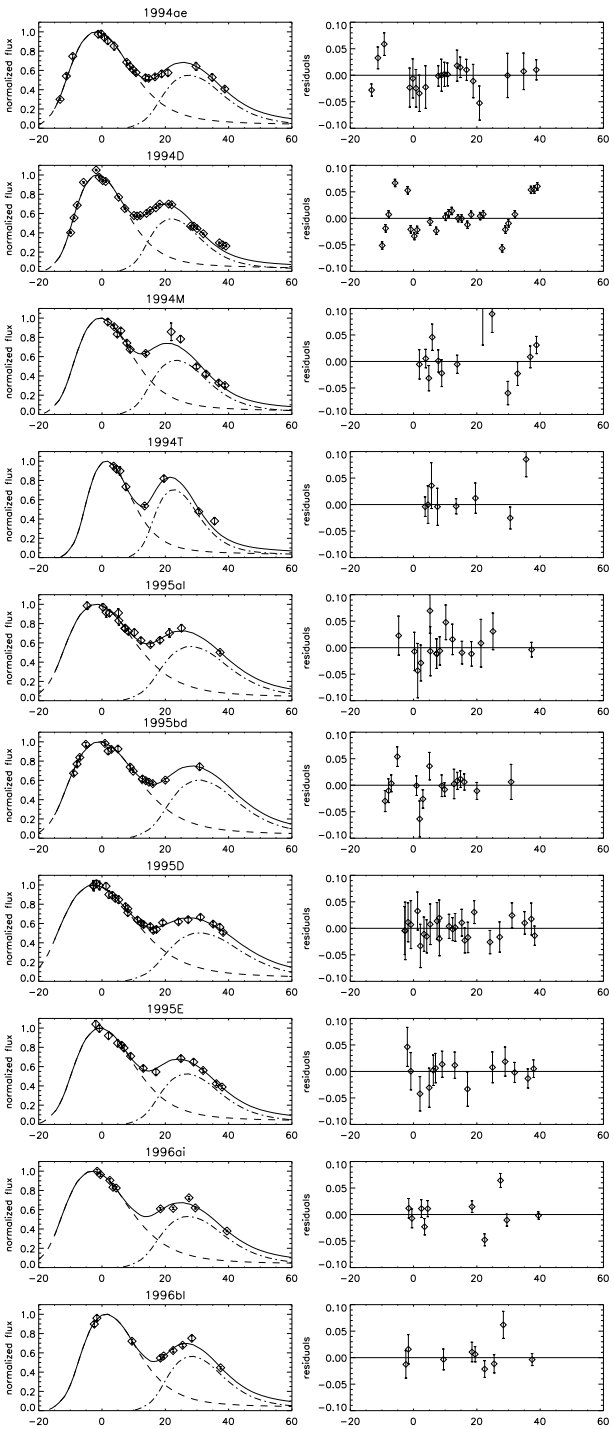


Fig. 1 continued. *I*-band light curve fits. On the ordinate is the the flux normalised to the first peak, on the abscissa the restframe time since *B*-band maximum. The dashed line and the dash-dotted line represent the two templates used to fit the first and second peak respectively.

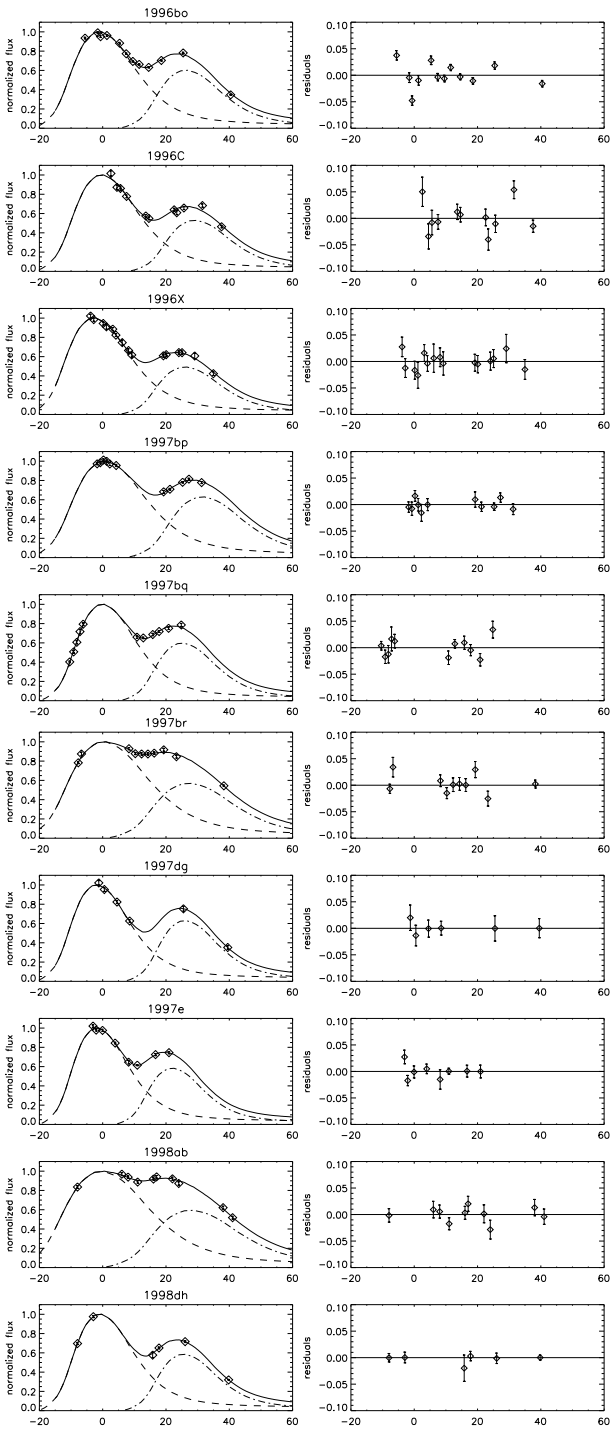


Fig. 1 continued. *I*-band light curve fits. On the ordinate is the flux normalised to the first peak, on the abscissa the restframe time since *B*-band maximum. The dashed line and the dash-dotted line represent the two templates used to fit the first and second peak respectively.

photometry. SN 1995E has been excluded from the sample as it is highly reddened (see also discussion in Nobili et al. (2003)). Two supernovae in the sample, the spectroscopically peculiar SN 1998es and SN 1999dq, plotted with filled symbols in figure 3, appear intrinsically redder than average, and become $\sim 2-3\sigma$ deviant from average after correction for host galaxy extinction. Before introducing light curve shape corrections, the spread measured in the $M^I_{\rm max}$ excluding these two SNe, is about 0.24 mag (0.28 mag if they are included). The solid line shows the best fit to the data, obtained for a slope $\alpha_I = 1.18 \pm 0.19$ and an absolute magnitude for a stretch $s_B = 1$

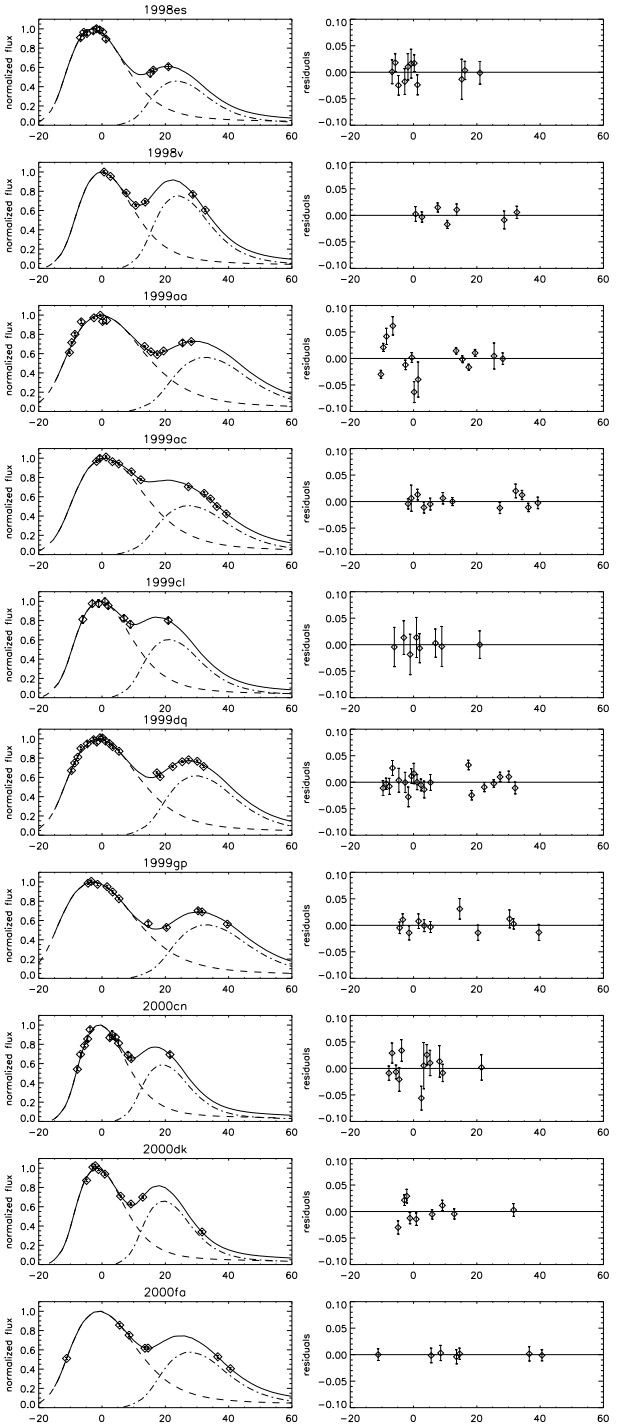


Fig. 1 continued. *I*-band light curve fits. On the ordinate is the flux normalised to the first peak, on the abscissa the restframe time since *B*-band maximum. The dashed line and the dash-dotted line represent the two templates used to fit the first and second peak respectively.

supernova equal to $M^I_{\rm max}(s_B = 1) = -18.84 \pm 0.03$ mag [2]. The dispersion measured as r.m.s. of the data about the fitted line is $0.17 \pm 0.03$ mag. A similar correlation was found between the peak magnitude and the stretch in *I*-band, $s_I$, with an

[2] The value fitted for $M^I_{\rm max}$ depends on the value assumed for the Hubble parameter, $H_0 = 72$ km s$^{-1}$ Mpc$^{-1}$. However, its value is not used in any of the further analysis presented in this paper.

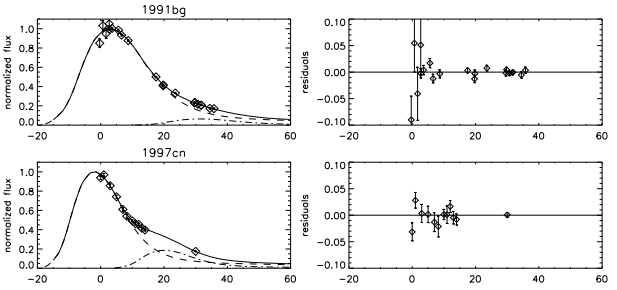


Fig. 1 continued. *I*-band light curve fits of the underluminous supernovae SN 1991bg and SN 1997cn. The dashed line and the dash-dotted line represent the two templates used to fit the first and second peak respectively. Note that the second peak is fitted as $\sim 3$ mag fainter than the first peak.

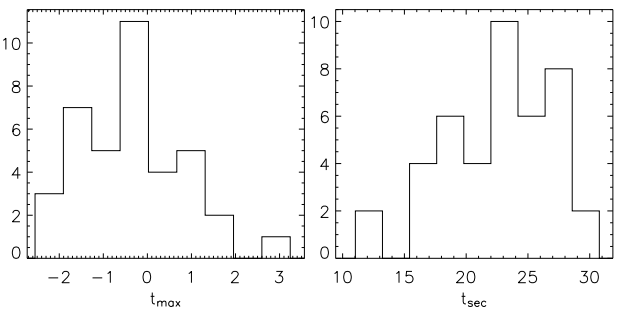


Fig. 2 Distribution of the time of *I*-band maximum referred to the time of *B*-band maximum (left panel) and the distribution of the time of second maximum referred to the time of *B*-band maximum (right panel).

r.m.s. of $\sim 0.19$ mag, about the best fit line, obtained excluding SN 1998es and SN 1999dq.

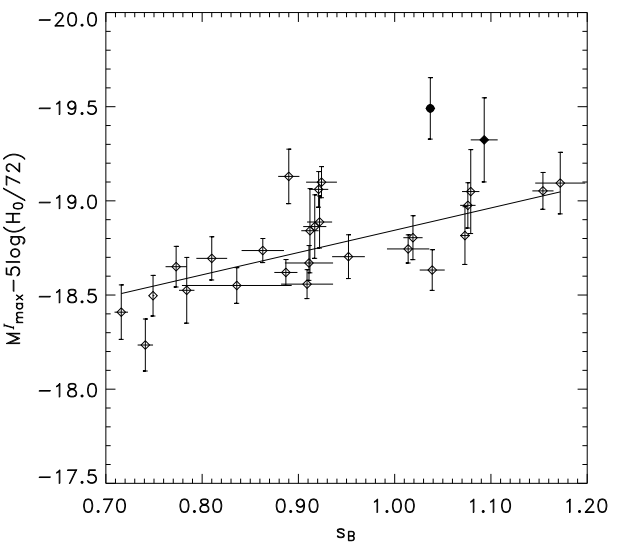


Fig. 3 *I*-band absolute magnitude versus stretch in B-band. The best fit gives $\alpha_I = 1.18 \pm 0.19$ and $M^I_{\rm max}(s_B = 1) = -18.84 \pm 0.03$ mag. The two deviating supernovae, SN 1998es (filled diamond) and SN 1999dq (filled circle), were excluded from the fit.

A correlation was found between $t_{\rm sec}$ and the *B*-band stretch factor, as shown in figure 4. There are three outliers labelled in

the figure, SN 1993H , SN 1998es and SN 1999ac, which are identified as spectroscopically peculiar supernovae. However, other supernovae in our sample also classified as spectroscopically peculiar, behave as "normal" type Ia SNe. We note that the *B*-band stretch factor for SN 1999ac is not well defined due to an asymmetry of the *B*-band light curve (Phillips et al., 2003). A linear fit to the data was performed, excluding the outlier SNe.

Figure 5 shows a possible correlation between $I_{\rm sec}$ and the stretch $s_B$, at least for $s_B < 0.9$, after correcting for the luminosity distance and both Galaxy and host galaxy extinction. This correlation, however, disappears for larger values of $s_B$.

All of these correlations, shown in figure 3 - 5, were expected since it has been suggested that the location and the intensity of the secondary peak depends on the *B*-band intrinsic luminosity of the supernova (Hamuy et al., 1996a).

**Figure 6 shows the *I*-band stretch, $s_I$, plotted versus *B*-band stretch, $s_B$, where some of the supernovae, three of which are spectroscopically peculiars, are two standard deviation away from the rest of the sample.**

We have investigated the possible existence of further relations between the fitted parameters, but find no additional statistically significant correlations.

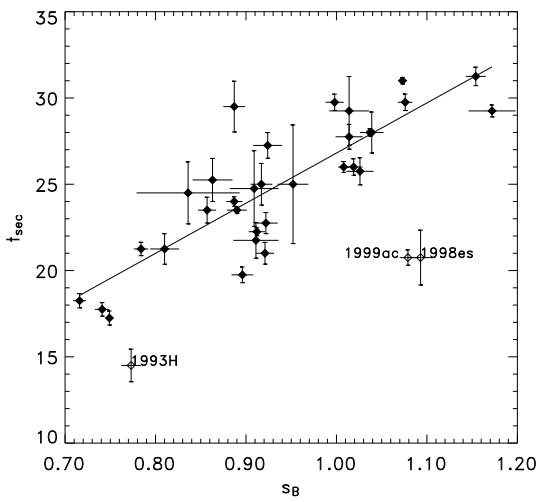


Fig. 4 Time of the second peak since $B_{\rm max}$ vs the stretch in *B*-band. The supernovae labelled are all classified as spectroscopically peculiar.

## 3. The *I*-band Hubble Diagram

The fitted values of $I_{\rm max}$ were used to build a Hubble diagram in the *I*-band. A total of 28 supernovae from the sample considered here have a redshift $z_{CMB} \geq 0.01$ for which peculiar velocities do not dominate the trend[3]. The maximum redshift in this sample is 0.1.

[3] The lower limit chosen in previous analyses by the SCP is slightly higher. However, we include these lower redshift SNe in the sample in order to increase the statistical significance. Cutting the Hubble diagram above $z = 0.015$ would decrease the sample by about 30 %. Note, however, that this choice does not significantly affect any of the results.

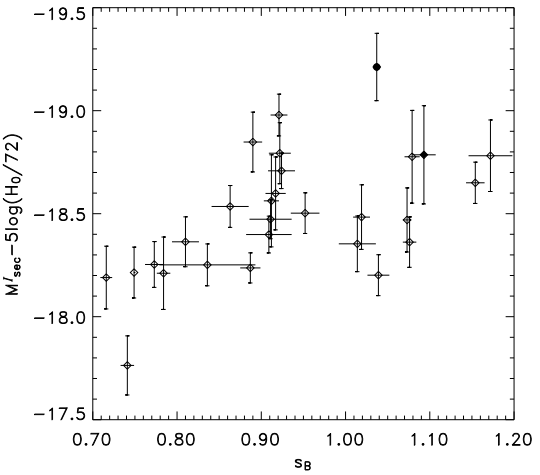


Fig. 5 Absolute magnitude of the secondary peak versus the stretch in *B*-band. The two deviating supernovae in figure 3 are SN 1998es (filled diamond) and SN 1999dq (filled circle).

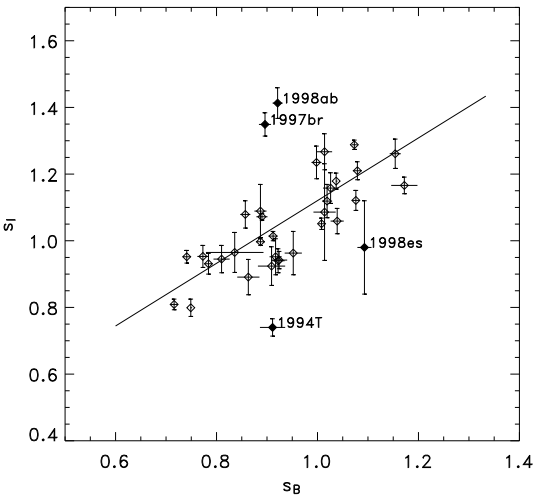


Fig. 6 *I*-band light curve stretch, $s_I$, versus *B*-band stretch, $s_B$. **The labelled supernovae are more than two standard deviation from the correlation shown by the ensemble. SN 1998ab, SN 1997br and SN 1998es are classified as peculiar.**

The width-luminosity relation found between the fitted absolute *I*-band magnitude and the *B*-band stretch factor was used to correct the peak magnitude, with a $\alpha_I = 1.18 \pm 0.19$ as measured in the previous section, similarly to what is usually done in the *B*-band (Perlmutter et al., 1999). The peak magnitude was also corrected for Milky Way and for host galaxy extinction:

$$m_I^{\rm eff} = m_I + \alpha_I(s_B - 1) - A_I^{host} - A_I^{MW} \quad (1)$$

The effective magnitude, $m_I^{\rm eff}$ of the nearby supernovae, listed in Table 3, have been used for building the Hubble diagram in *I*-band, shown in figure 7. The inner error bars include an uncertainty in the redshifts due to peculiar velocities of the host galaxies, assumed to be 300 km s$^{-1}$.

| SN | $s_B$ | $z_{CMB}$ | $m_I^{\rm eff}$ |
|---|---|---|---|
| 1992al | 0.917 ± 0.012 | 0.014 | 14.868 ± 0.162 |
| 1992bc | 1.076 ± 0.008 | 0.020 | 15.718 ± 0.111 |
| 1992bg | 0.952 ± 0.017 | 0.036 | 17.120 ± 0.107 |
| 1992bh | 1.014 ± 0.022 | 0.045 | 17.637 ± 0.062 |
| 1992bo | 0.741 ± 0.008 | 0.017 | 15.710 ± 0.138 |
| 1992bp | 0.863 ± 0.022 | 0.079 | 18.689 ± 0.055 |
| 1993H | 0.773 ± 0.011 | 0.025 | 16.169 ± 0.106 |
| 1993O | 0.887 ± 0.012 | 0.053 | 17.967 ± 0.054 |
| 1993ag | 0.909 ± 0.027 | 0.050 | 17.928 ± 0.069 |
| 1994M | 0.810 ± 0.016 | 0.024 | 16.080 ± 0.111 |
| 1994T | 0.911 ± 0.025 | 0.036 | 17.105 ± 0.085 |
| 1995bd | 1.172 ± 0.026 | 0.014 | 14.939 ± 0.162 |
| 1996C | 1.039 ± 0.013 | 0.027 | 16.669 ± 0.097 |
| 1996bl | 0.924 ± 0.016 | 0.035 | 16.630 ± 0.070 |
| 1996bo | 0.890 ± 0.011 | 0.016 | 14.860 ± 0.138 |
| 1997bq | 0.912 ± 0.009 | 0.010 | 14.154 ± 0.219 |
| 1997dg | 0.836 ± 0.057 | 0.030 | 16.740 ± 0.110 |
| 1997E | 0.784 ± 0.008 | 0.013 | 14.888 ± 0.173 |
| 1998ab | 0.921 ± 0.010 | 0.028 | 16.180 ± 0.083 |
| 1998es | 1.093 ± 0.014 | 0.010 | 13.885 ± 0.219 |
| 1998V | 0.922 ± 0.013 | 0.017 | 15.272 ± 0.131 |
| 1999aa | 1.073 ± 0.005 | 0.015 | 15.250 ± 0.146 |
| 1999ac | 1.079 ± 0.009 | 0.010 | 14.144 ± 0.218 |
| 1999dq | 1.037 ± 0.000 | 0.014 | 14.382 ± 0.155 |
| 1999gp | 1.154 ± 0.011 | 0.026 | 16.303 ± 0.090 |
| 2000cn | 0.749 ± 0.000 | 0.023 | 16.114 ± 0.107 |
| 2000dk | 0.716 ± 0.007 | 0.016 | 15.375 ± 0.147 |
| 2000fa | 1.019 ± 0.010 | 0.022 | 16.029 ± 0.106 |

Table 3 List of SNe used in the Hubble diagram. $m_I^{\rm eff}$ is the peak magnitude corrected for the dust extinction and for the width-luminosity relation, following Eq. 1. The quoted uncertainties do not include the redshift contribution due to peculiar velocities in the host galaxies, assumed equal to 300 km s$^{-1}$.**Redshifts from Table 2 are transformed in CMB frame.**

The solid line represents the best fit to the data for the concordance model $\Omega_M = 0.25$ and $\Omega_\Lambda = 0.75$. The fitted parameter, $\mathcal{M}_I$, is defined (as in Perlmutter et al. (1997)) to be

$$\mathcal{M}_I \equiv M_I - 5 \log H_0 + 25 \qquad (2)$$

where $M_I$ is the $I$-band absolute magnitude for a $B$-band stretch $s_B = 1$ supernova. The two redder supernovae, SN 1998es and SN 1999dq, were excluded from the fit, and are plotted with different symbols in figure 7. The value fitted is $\mathcal{M}_I = -3.19 \pm 0.03$. In order to disentangle the intrinsic dispersion from the statistical scatter due to the measurement uncertainties, we simulated data sets with a dispersion given by the measurement uncertainty only. Since the uncertainty due to peculiar motion of the host galaxy is dominant at very low redshift ( ∼0.2 mag for z =0.01), we limited this calculation to only 15 SNe with z > 0.025, which correspond to a peculiar velocity uncertainty of the same order as the measurement uncertainties in our sample. The average of the r.m.s. measured on each of the simulated data sets is geometrically subtracted from the dispersion measured as r.m.s. on the data (0.17 mag), resulting in $\sigma = 0.13$ mag. We consider this an estimate of the intrinsic dispersion of the stretch corrected $I$-band light curve maximum, which agrees with the estimate given by Hamuy et al. (1996b) using 26 SNe of the Calan/Tololo sample. The estimated intrinsic uncertainty of 0.13 mag has been added in quadrature to the outer error bars of the plotted data. Note that if no correction $\alpha_I(s_B - 1)$ is applied the dispersion in the Hubble diagram becomes 0.24 ± 0.04 mag, somewhat smaller than the corresponding dispersion measured in the "uncorrected" $B$-band Hubble diagram. Moreover, we computed the dispersion in the Hubble diagram for the three data sets separately. **The small differences found are compatible with statistical fluctuations.**

## 4. High redshift supernovae

Next, we explore the possibility of extending the Hubble diagram to higher redshifts, where the effects of the energy density components of the universe are, in principle, measurable. The restframe $I$-band data available to date for this purpose are unfortunately very limited. It consists of only three supernovae at redshift $z \sim 0.5$ observed in the near infrared (NIR) J-band collected during three different campaigns conducted using different facilities and by two different teams. Keeping all of these sources of systematic errors in mind, we include the three supernovae in the Hubble diagram to show its potential and complementarity with respect to the standard $B$-band Hubble diagram.

The following sections describe the three high redshift supernovae, SN 2000fr, SN 1999ff and SN 1999Q, and the light curve fitting procedure used.

### 4.1. SN 2000fr

SN 2000fr was discovered by the Supernova Cosmology Project (SCP) during a search for type Ia supernovae at redshift $z \sim 1$ conducted in $I$-band with the CFH-12k camera on the Canada-France-Hawaii Telescope (CFHT)(Schahmaneche et al., 2001). The depth of the search allowed us to discover this $z \sim 0.5$ supernova during its rise, about 11 rest-frame days before maximum B-band light.

The supernova type was confirmed with two spectra taken at the Keck II telescope and VLT, showing that it was a normal type Ia at $z = 0.543$, see Lidman et al. (2004); Garavini et al. (2004a) for an extensive analysis of the spectrum. This supernova was followed in restframe $B$, $V$ and $I$ filters involving both ground and space based facilities. Approximately one year later, when SN 2000fr had faded sufficiently, infrared and optical images of the host galaxy were obtained. The optical light curves in (Knop et al., 2003) were re-fitted using the improved spectral templates. We found a stretch parameter $s_B = 1.034 \pm 0.013$ and a time of $B_{\rm max}$, $t_{\rm max}$= MJD 51685.6. Restframe $B-V$ measurements at the time of $B_{\rm max}$ indicate that SN 2000fr did not suffer from reddening due to dust in the host galaxy, (see also Section 6 for a more extensive discussion). The adopted Milky Way reddening is $E(B-V) = 0.030$ mag (Schlegel et al., 1998).

The near-infrared data were collected with ISAAC at the VLT. They consist of $J_s$-band observations during three epochs and a final image of the host galaxy without the SN (see Table

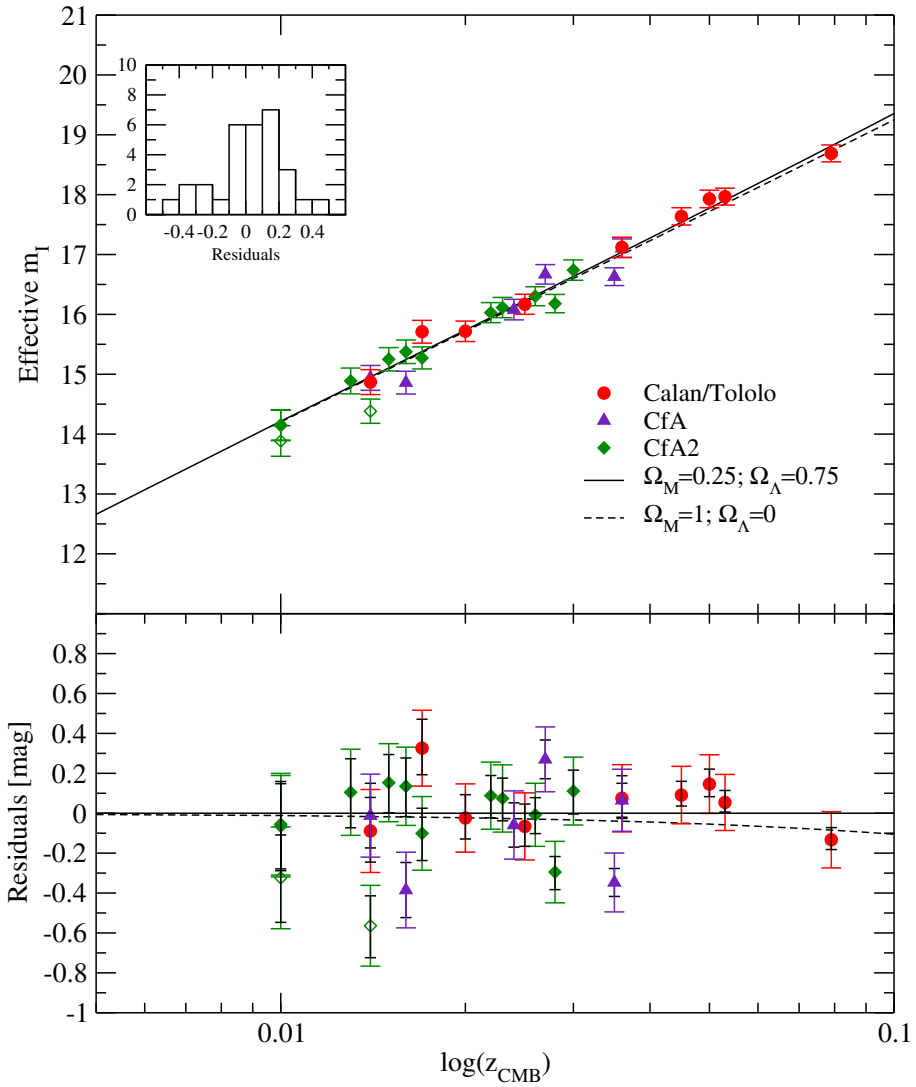


Fig. 7 Effective *I*-band maximum vs redshift for the nearby supernovae of the Calan/Tololo, CfA and CfA2 sample. The data have been corrected for the stretch-luminosity relation and for Milky Way and host galaxy extinction. The r.m.s. along the concordance model line is $\sigma = 0.17 \pm 0.04$ mag. Subtracting the contribution of the average uncertainty, results in 0.13 mag estimated intrinsic dispersion (see text for details). SN 1998es and SN 1999dq were excluded from the fit (see text) and are plotted with open diamonds. The inset plots the histogram of the residuals.

4). Each data point is composed of a series of 20 to 60 images with random offsets between exposures. Figure 8 shows a comparison between the *J* Persson filters and the narrower $J_s$ filter used for the observations, together with the atmospheric transmission, and the spectral template at maximum.

The advantage in using the narrower $J_s$ filter is that the transmission of the filter is not determined by the region of strong atmospheric absorption between 13500 and 15000 Å. Consequently, the zero-point is significantly more stable than that of standard *J*. This was very useful, because all of the ISAAC data were taken in queue mode, where typically only one or two standard stars, chosen from the list of Persson et al. (1998), are observed during a night. All data, except the reference images, were taken during photometric nights and the difference in the zero-points from one night to the next was less than 0.01 magnitudes.

| MJD | Epoch | $J_s$(mag) | *I*(mag) |
|---|---|---|---|
| 51685.06 | -0.33 | 22.50 ± 0.09 | 23.52 ± 0.10 |
| 51709.02 | 15.20 | 23.57 ± 0.22 | 24.52 ± 0.23 |
| 51731.96 | 30.07 | 23.14 ± 0.15 | 23.99 ± 0.16 |

Table 4 Summary of $J_s$ data for SN 2000fr. The quoted errors are due to statistical Poisson noise and the uncertainty on the ZP (contributing 0.01 mag). Epochs are in restframe relative to *B*-band maximum. Restframe *I*-band is obtained through cross-filter *K*-correction from the observed $J_s$-band to Bessel *I*-band, and addition of the offset found between optical and IR systems. The uncertainties also include the contribution from *K*-corrections, estimated to be 0.05 mag at all epochs considered.

The data were reduced using both internally developed routines and the XDIMSUM package in IRAF[4]. We found system-

[4] IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities

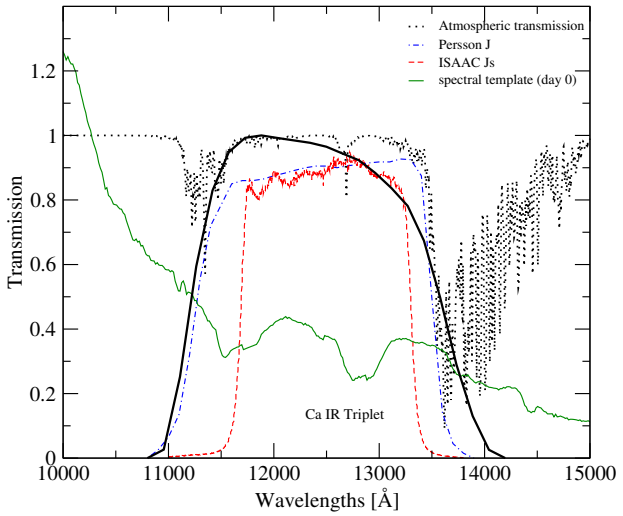


Fig. 8 Comparison between the standard $J$ Persson filter (dashed-dotted line), the $J_s$ filter at ISAAC used for the observations of SN 2000fr (dashed line) and the $I$-band red-shifted to z=0.543 (heavy solid line). The atmospheric transmission is also plotted (dotted line). The spectral template at day 0 is on an arbitrary flux scale for readability purpose (solid line).

atic differences between the two analyses, which are within the quoted uncertainties. The supernova images were aligned with the host galaxy images and the flux scaled to the one with best seeing, using the field stars before performing PSF photometry. The results are presented in Table 4. The stated uncertainties include the statistical Poisson noise and the uncertainty on the estimate of the zero point, added in quadrature.

The $J_s$-band magnitude takes into account a colour term which arises from the difference between the $J$ filter of the standard star system and the $J_s$ filter used in ISAAC. This correction was small $\sim 0.012$ mag.

The cross-filter $K$-correction, $K_{IJ_s}$, to convert from $J_s$-band to rest-frame $I$-band, has been calculated following Kim et al. (1996) using the spectral templates improved for this work. A term has been added to account for the appropriate transformation between IR and optical photometric systems, equal to $(I - J) = 0.03(\pm 0.02)$, determined by using the Vega magnitudes in $I$ and $J$ (Bessell et al., 1998; Cohen et al., 2003). We conservatively assume 0.05 mag total uncertainty in the $K$-corrections (see Section 2.2).

### 4.2. SN 1999ff

SN 1999ff was discovered by the High-Z Supernova Search Team (HZSST) during a search conducted at CFHT using the CFH-12k camera in $I$-band (Tonry et al., 2003).[5] The supernova was confirmed spectroscopically as a type Ia at redshift $z = 0.455$. The adopted Milky Way reddening is $E(B - V) = 0.025$ mag (Schlegel et al., 1998).

$J$-band observations, corresponding to restframe $I$-band, reported in (Tonry et al., 2003), were taken at Keck using NIRC at two epochs only. The $J$-band filter available at Keck is very similar to the ISAAC-$J_s$, shown in figure 8. We have used the published observed photometry, and, for consistency with the treatment of both the low redshift supernovae and SN 2000fr, we computed the $K$-corrections using the improved spectral templates. We found differences with the results published in Tonry et al. (2003), due to the use of an incorrect filter in the originally published results, however, the $K$-corrections calculated as part of the MLCS distance fits to this object were done with the correct filter (Brian Schmidt, private communication). The $I$-band magnitudes were also corrected for the offset found between the optical and IR systems, as explained in the previous section. The restframe $I$-band magnitudes obtained this way are reported in Table 5. The published optical $R$-band data were used to fit the restframe $B$-band light curve using the stretch method. The time of maximum was within 1 day of the Tonry et al. (2003) value, with a best fit for the stretch $s_B = 0.80 \pm 0.05$.

| MJD | Epoch | I(mag) |
|---|---|---|
| 51501.29 | 5.01 | 23.57 ± 0.11 |
| 51526.31 | 22.21 | 24.06 ± 0.24 |

Table 5 Summary of IR data for SN 1999ff. Epochs are in restframe relative to $B$-band maximum ($t_{\max} = MJD51494.8$); restframe $I$-band magnitudes are computed applying $K$-corrections to the observed $J$-band data published in Tonry et al. (2003), and adding the offset found between optical and IR systems. The uncertainties also include the contribution from $K$-corrections, estimated to be 0.05 mag at all epochs considered.

### 4.3. SN 1999Q

SN 1999Q was discovered by the HZSST using the CTIO 4m Blanco Telescope and was spectrally confirmed to be a Type Ia SN at $z = 0.46$ Garnavich et al. (1999). The adopted Milky Way reddening is $E(B-V) = 0.021$ magnitudes(Schlegel et al., 1998).

SN 1999Q was observed in the $J$-band over five epochs, the first with SofI on the ESO NTT and the following four with NIRC at Keck Telescope (Riess et al., 2000). **We have re-analysed the publicly available SofI data and found a $J$-band measurements, $J$=22.63 ± 0.15 mag, brighter than the value published by Riess et al. (2000),** $23.00 \pm 0.14$ **mag. Due to this large discrepancy, corresponding to about two standard deviations, we decided to discard this SN from the rest of the analysis.**

### 4.4. Light Curve fits for the high redshift supernovae.

The $I$-band light curves of the high redshift supernovae are not as well sampled in time as the low redshift sample analysed in

for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

[5] Another supernova, SN 1999fn, was followed in $J$-band by the HZSST during the same search. However since it was found in a highly extincted Galactic field, E(B-V)=0.32 mag, and since it was strongly contaminated by the host galaxy, we did not include it in our analysis.

Section 2. There are only few data points for each SN, making it impossible to perform the 5 parameter fit. Thus, we used the results of the fit of the local sample of supernovae to build a set of 42 *I*-band templates, which in turn have been used to fit the high redshift SN light curves.

The best fit light curve for each of the 42 supernovae in our low-redshift sample can be viewed as defining an *I*-band template. The high-redshift supernovae are fit to each template with a single free parameter, $I_{\max}$, the absolute normalisation of the template. The time of $B_{\max}$ is obtained from the literature (SN 1999ff) or from our own *B*-band light curve fits (SN 2000fr). The best-fitting low-redshift *I*-band template fixes the date of the I-band maximum relative to the date of the *B*-band maximum. A $\chi^2$ comparison was used to choose the best low redshift template.

Figs. 9 and 10 show the comparison of the data with the best fit template for each of the supernovae. Table 6 gives the results of the fit together with redshift, the number of data points, the template giving the best fit and the $\chi^2$. As there are few data points for each SN, the $\chi^2$ parameter has little significance for estimating the goodness of the fits. Thus, to estimate the possible systematic error in the measured peak magnitude from the selection of light curve template, we computed the r.m.s. of the fitted $I_{max}$ of all the light curve templates giving a $\chi^2 \leq \chi^2_{min} + 3$. This systematic uncertainty is reported also in Table 6. In the cases of SN 1999ff and SN 2000fr this is quite small, and compatible with the scatter due to the statistical uncertainties, thus, it is a conservative estimate.

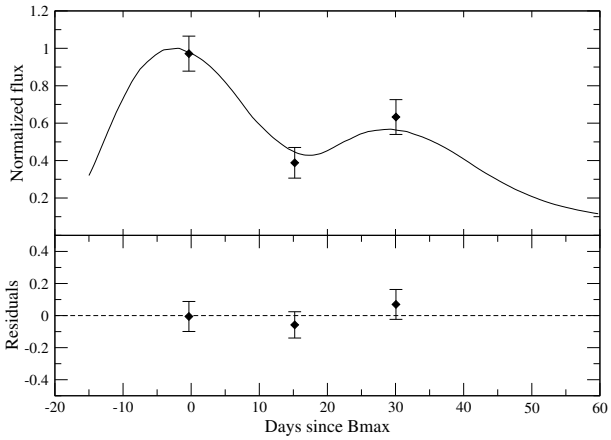


Fig. 9 *I*-band fit for SN 2000fr. Best fit was obtained with the template of the nearby SN 1992bc, out of 42 templates. The fit was performed with only one free parameter, the peak magnitude, $I_{\max} = 23.48 \pm 0.08$ mag. Supplemental data from the *B*-band (not shown) is used to fix the date of *B*-maximum.

### *4.5. Monte-Carlo test of the fitting method*

A Monte-Carlo simulation was run in order to test the fitting method applied to the high redshift SNe. The measurement uncertainties were used to generate a set of 1000 SNe, with data points randomly distributed around the real data and at the same epochs as the data. All the simulated data sets were in turn fitted with the 42 templates and the one giving the minimum $\chi^2$ was selected for each simulation. The distribution of the fitted parameters in each of the simulated data sets around the true values, fitted on the experimental data, was studied to check for systematic uncertainty in the fitting procedure. This was found to be robust, always selecting the same template as the one giving the best fit for all the three SNe. No bias was found, therefore confirming the peak magnitude fitted with this method. The uncertainty on $I_{\max}$ reported in Table 6 was consistent with the dispersion in the distribution measured from the simulations.

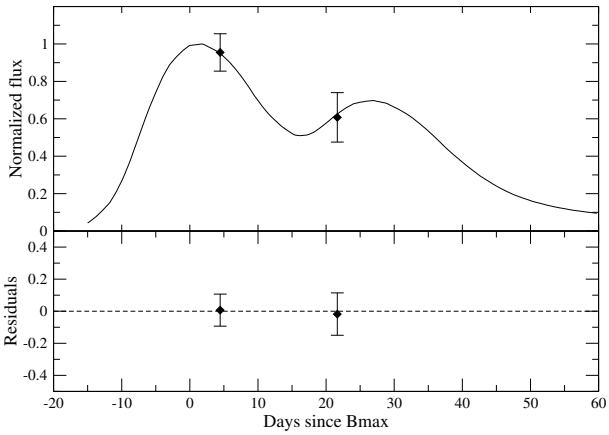


Fig. 10 *I*-band fit for SN 1999ff. Best fit was obtained with the template of SN 1999bl, out of 42 templates. The fit was performed with only one free parameter, the peak magnitude, $I_{\max} = 23.55 \pm 0.10$ mag. Supplemental data from the *B*-band (not shown) is used to fix the date of *B*-maximum.

## 5. The *I*-band Hubble diagram up to $z \sim 0.5$

The *I*-band peak magnitudes of the high redshift supernovae reported in Table 6 were corrected for Milky Way extinction. Note that all SNe have been reported not to suffer from extinction in their host galaxies.

The Hubble diagram has been built both without and with width-luminosity correction (case *a* and case *c* respectively), where the systematic uncertainties on the peak magnitudes of the distant supernovae, listed in Table 6, are added in quadrature to the statistical uncertainties. Cases *b* and *d* are like *a* and *c* but neglect the systematic uncertainties from Table 6. Figure 11 shows the extended Hubble diagram (case *c*), where an intrinsic uncertainty of 0.13 mag has been added in quadrature to the measurement errors of the plotted data. The solid line represent the best fit to the nearby data for the concordance model $\Omega_M = 0.25$ and $\Omega_\Lambda = 0.75$. Also plotted is the model for $\Omega_M = 1$ and $\Omega_\Lambda = 0$ (dashed line), and a flat, $\Lambda = 0$ universe in presence of a homogeneous population of large grain dust in the intergalactic (IG) medium able to explain the observed dimming of type Ia SNe at $z \sim 0.5$ in the *B*-band (dotted line) (Aguirre, 1999a,b). The bottom panel shows the residuals obtained for case *a*. Table 7 lists the $\chi^2$ values for the high redshift SNe for each of the models. The low statistics of the high redshift sample is insufficient to draw strong conclusions on cosmological parameters.

| SN | $z$ | $s_B$ | $n$ | $I_{\max}$ | template | $\chi^2$ | $A_J^{MW}$ |
|---|---|---|---|---|---|---|---|
| SN 2000fr | 0.543 | 1.034 ± 0.011 | 3 | 23.48 ± 0.08 ± 0.04 | SN 1992bc | 1.04 | 0.027 |
| SN 1999ff | 0.455 | 0.80 ± 0.05 | 2 | 23.55 ± 0.10 ± 0.08 | SN 1996bl | 0.05 | 0.022 |

Table 6 List of the high redshift type Ia SNe used in this work. Columns are: redshift, number of data points used in the fit, magnitude of the peak resulted from the fit (both statistical and systematic uncertainties are given) before Milky Way extinction correction, best fit template, $\chi^2$ of the fit, Milky Way extinction in the $J$-band.

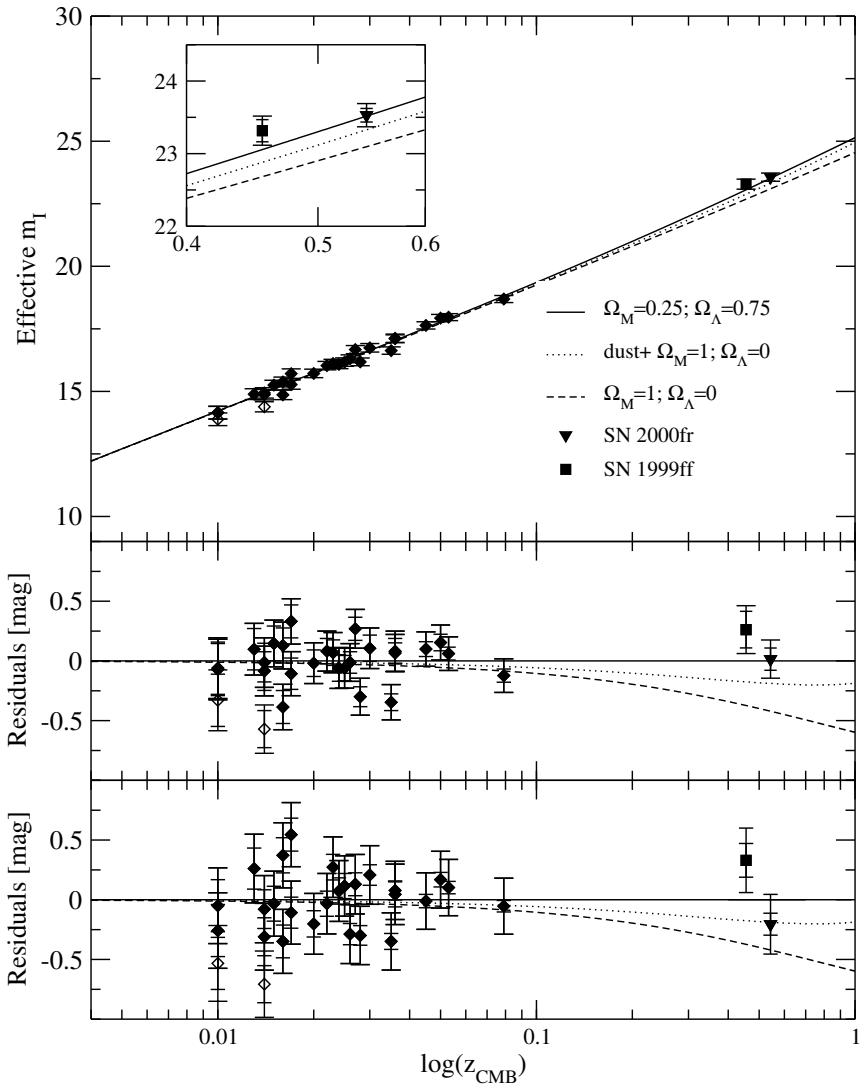


Fig. 11 Effective $I$-band maximum vs redshift for the nearby supernovae of the Calan/Tololo, CfA and CfA2 sample, together with three supernovae at redshift $\sim 0.5$ for case $c$ (top panel), residuals to the $(\Omega_M, \Omega_\Lambda)$=(0.25,0.75) model for case $c$ (middle panel) and case $a$ (bottom panel). SN 1998es and SN 1999dq were excluded from the fit (see text) and are plotted with open diamonds.

Although these results are encouraging, additional unidentified systematic uncertainties on the distant SNe cannot be excluded at this point.

Some uncertainties are specific to the sample considered here. The different fitting methods applied to the restframe $I$-band light curve for the low and high redshift samples can be easily overcome if distant supernovae are followed at NIR wavelengths with better time coverage. Both the low and high redshift samples used in this analysis are rather heterogeneous, as they were collected from different data sets. Future data sets collected with a single instrument would naturally solve this problem.

## 6. SN colours and study of intergalactic dust

Multi-colour photometry allows for extinction tests for non-standard dust having only weak wavelength dependence, such as a homogeneous population of large grain dust, as proposed by Aguirre (1999a,b). Assuming a density of "grey" dust in the IG medium to explain the observed dimming of supernovae at redshift $z = 0.5$ in $B$-band with $\Omega_M = 1$ and $\Omega_\Lambda = 0$, we

| $(\Omega_M, \Omega_\Lambda)$ | $\chi^2_a$ | $\chi^2_b$ | $\chi^2_c$ | $\chi^2_d$ |
|---|---|---|---|---|
| (0.25,0.75) | 1.09 | 0.72 | 0.58 | 1.18 |
| (1,0) | 3.78 | 9.15 | 7.78 | 4.22 |
| $(1,0)_{dust}$ | 1.76 | 3.24 | 2.68 | 1.98 |

Table 7 Reduced $\chi^2$ (for 2 *dof*) of the high redshift data to each model, without stretch correction and with systematic uncertainties added in quadrature ($\chi^2_a$), neglecting the systematic uncertainties ($\chi^2_b$), with stretch correction and adding the systematic uncertainties in quadrature ($\chi^2_c$) or neglecting them ($\chi^2_d$).

calculate the expected colour extinction following Goobar et al. (2002a) using the SNOC Monte-Carlo package (Goobar et al., 2002b), for two cases of $R_V = 4.5$ and $R_V = 9.5$ and a constant comoving dust density between z=0 and the SN redshifts.

For both SN 1999ff and SN 2000fr we computed rest frame $B-V$ and $B-I$. Table 8 and 9 lists the colours for all the SNe, corrected for Milky Way extinction. The colour evolution has been compared to models for a $\Lambda$ dominated universe and a $\Omega_M = 1$, $\Omega_\Lambda = 0$ universe in presence of grey dust ($R_V$=4.5 and $R_V$=9.5) accounting for the faintness of type Ia supernovae at $z \simeq 0.5$. This is shown in figure 12. The error bars include the intrinsic colour dispersion contribution.

The $\chi^2$ has been computed for both $B-V$ and $B-I$ evolution for SN 1999ff and SN 2000fr, and for both supernovae together (see Table 10). The correlations between SN colours at different epochs found in (Nobili et al., 2003) were taken into account. However, we note that, although this correlation should be taken into account in the calculations, neglecting it would not change significantly the conclusions of the analysis. Although individual supernovae give $\chi^2$ values that would seem to distinguish between the models being compared, their combination disfavour such conclusions.

| Epoch | $B-I$ |
|---|---|
| SN 2000fr | |
| -0.32 | -0.51 ± 0.12 |
| 14.70 | -0.50 ± 0.24 |
| 29.08 | 1.42 ± 0.17 |
| SN 1999ff | |
| 5.59 | -0.18 ± 0.13 |
| 27.20 | 1.38 ± 0.24 |

Table 8 Restframe $B-I$ colours in magnitudes for the three high redshift SNe. Epoch is in days relative to $B$-band maximum, divided by the $B$-band stretch.

To make our test for grey dust more robust a different approach was followed. The method of least squares has been used to combine colour measurements along time for each supernova (see Cowan, 1998, p.106 for details). The residuals of each supernova colour from the expected model, plotted in figure 12, are weight-averaged together, and the covariance matrix is used as weight in the calculation. In the following, we refer to $E(X-Y)$, to describe the colour excess of any supernova with respect to the $X-Y$ average colour of local supernovae, as derived in (Nobili et al., 2003). First we applied this method

| Epoch | $B-V$ |
|---|---|
| SN 2000fr | |
| -7.97 | -0.06 ± 0.05[a] |
| -3.51 | -0.14 ± 0.05[a] |
| 4.60 | -0.12 ± 0.05 |
| 12.93 | 0.24 ± 0.08 |
| 20.31 | 0.61 ± 0.07 |
| 30.22 | 0.99 ± 0.09 |
| SN 1999ff | |
| -7.99 | 0.03 ± 0.08[a] |
| 1.91 | -0.02 ± 0.09 |
| 1.98 | 0.10 ± 0.12 |
| 2.91 | 0.23 ± 0.12 |
| 19.55 | 0.71 ± 0.09 |
| 28.75 | 1.22 ± 0.20 |

Table 9 Restframe $B-V$ colours in magnitudes for the two of the high redshift SNe. Epoch is in days relative to $B$-band maximum, divided by the $B$-band stretch.

[a] data not included in the analysis because out of the range in which Nobili et al. (2003) studies colour correlations.

to all local supernovae and used the results to establish the expected distribution in the $E(B-I)$ vs $E(B-V)$ plane, as showed in figure 13.

As the high redshift SNe were not corrected for host galaxy extinction, we computed the local sample distribution for two cases: the left panels represent the distribution of colour excess of 27 nearby SNe not corrected (top panel) and corrected (bottom panel) for host galaxy extinction, where the spectroscopically peculiar SNe have been excluded from the analysis. The projection of the ellipses on each colour axis is the estimated standard deviation of one colour distribution and the inclination is defined by the linear Pearson correlation coefficient computed on the same data sample. The solid contours represent 68.3%, 95.5% and 99.7% probability.

The right panels in figure 13 show the combined values of colour excess for the high redshift supernovae. These are compared to the local supernova distribution (solid lines), that represent the distribution expected in the absence of IG dust. Also plotted is the 68.3% level of the expected distribution in presence of grey dust with $R_V = 9.5$, represented by the ellipse (dashed line) displaced by (0.06,0.19) from the no-dust model. Only the case of $R_V = 9.5$ has been plotted for readability reasons, given the small difference between the two dust models. Note that this is the closer to the no-dust model. The ellipse corresponding to $R_V = 4.5$ would be displaced by (0.03,0.04), respectively in $E(B-V)$ and $E(B-I)$, from the latter.

We computed the $\chi^2$ of the high redshift data for each model (with and without IG dust), for both the situation in the top and bottom panels of figure 13. For each model, we sum the $\chi^2$ contribution from all SNe, taking into account the correlation found between $E(B-V)$ and $E(B-I)$ in the nearby sample. The reduced $\chi^2$ (for 4 degrees of freedom) are 0.63, 1.20, and 0.97 for the no-dust, IG dust with $R_V = 4.5$ and IG dust with $R_V = 9.5$ model respectively, for the host-galaxy extinction cor-

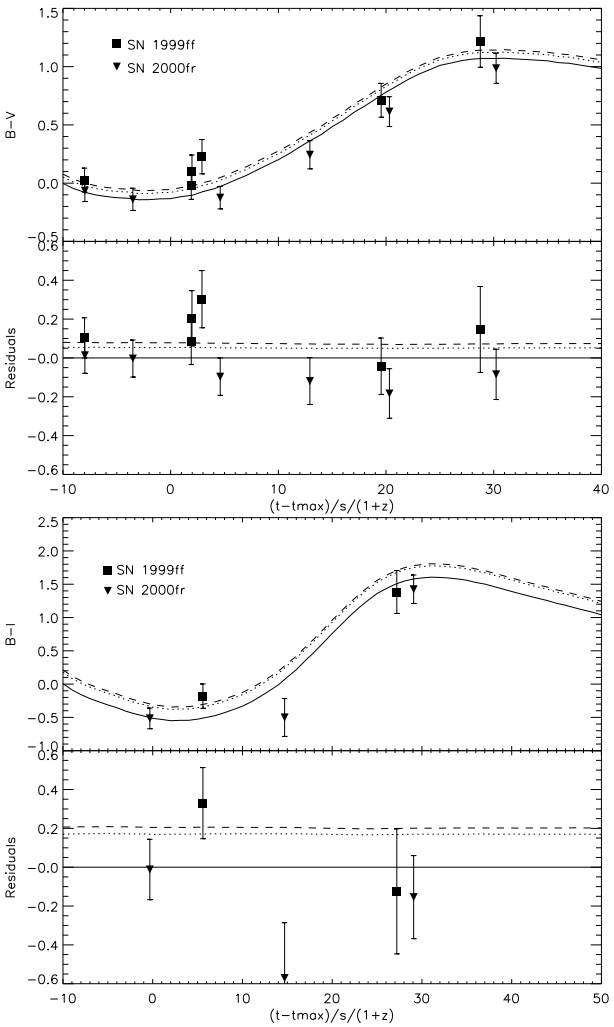


Fig. 12 High redshift SNe colour evolution, $B-V$ (top panel) and $B-I$ (bottom panel), compared to a $\Lambda$ dominated universe (solid line) and to an $\Omega_M = 1$, $\Omega_\Lambda = 0$ universe with presence of grey dust with $R_V$=4.5 (dashed line) and $R_V$=9.5 (dotted line).

rected case, and 0.62, 1.55 and 1.32 uncorrected. We note that both the intrinsic dispersion of the nearby data, represented by the ellipses in figure 13, and the uncertainties on the high redshift SNe, have been taken into account for computing the $\chi^2$. The statistical significance of these results is very limited, and should only be taken as an example of the method developed here.

A Monte Carlo simulation was used to estimate the minimum sample size needed to test for presence of homogeneously distributed grey dust in the IGM. SNe colours were generated following the binormal distribution defined by the SN local sample. Under the assumption that the systematic effects are negligible and an average measurement uncertainty of 0.05 mag in both $E(B-V)$ and $E(B-I)$, we found that a sample of at least 20 SNe would be needed to be able to exclude the IG dust model with $R_V = 9.5$ at the 95% C.L.

However, the possibility for this analysis to be affected by systematic effects is not negligible. Increasing the sample and the time sampling for each object would allow us not only to improve the significance of our statistic, but it will also be a means to identify and quantify systematic effects involved.

| | $\chi^2_{B-V}$/dof | $\chi^2_{B-I}$/dof |
|---|---|---|
| SN 2000fr | | |
| no dust, $(\Omega_M, \Omega_\Lambda)$=(0.25,0.75) | 2.33/4 | 2.08/3 |
| dust $R_v$ = 9.5, $(\Omega_M, \Omega_\Lambda)$=(1,0) | 4.24/4 | 3.94/3 |
| dust $R_v$ = 4.5, $(\Omega_M, \Omega_\Lambda)$=(1,0) | 5.29/4 | 4.50/3 |
| SN 1999ff | | |
| no dust , $(\Omega_M, \Omega_\Lambda)$=(0.25,0.75) | 6.05/5 | 3.93/2 |
| dust $R_v$ = 9.5, $(\Omega_M, \Omega_\Lambda)$=(1,0) | 4.69/5 | 2.06/2 |
| dust $R_v$ = 4.5, $(\Omega_M, \Omega_\Lambda)$=(1,0) | 4.31/5 | 1.89/2 |
| SNe combined | | |
| no dust , $(\Omega_M, \Omega_\Lambda)$=(0.25,0.75) | 8.38/9 | 6.01/5 |
| dust $R_v$ = 9.5, $(\Omega_M, \Omega_\Lambda)$=(1,0) | 8.93/9 | 6.00/5 |
| dust $R_v$ = 4.5, $(\Omega_M, \Omega_\Lambda)$=(1,0) | 9.59/9 | 6.39/5 |

Table 10 $\chi^2$ computed for the 3 different models and colours for each of the supernovae and for their combination.

## 7. Test for SN brightness evolution

Evolution of the properties of the supernova progenitors with redshift has been often proposed as an alternative explanation to the observed dimming of distant SNe. This is based on the assumption that older galaxies show different composition distribution than younger ones, e.g. an increased average metallicity, therefore offering different environmental conditions to the exploding star. A simple way to test for evolution is to compare properties of nearby SNe with distant ones. This will not prove that there is no evolution, but it will exclude it on a supernova-by-supernova or property-by-property basis, finding always counterparts of distant events in the local sample.

In this work we compared the colours of nearby and distant supernovae (primarily to test presence of "grey" dust). Although the size of the high redshift sample is very limited, our attempt does not give evidence for evolution of the average SN colours. Furthermore, the correlation found between the intensity of the secondary peak of *I*-band light curve and the supernova luminosity give an independent way of testing for evolution. The restframe *I*-band light curve for the high-redshift supernovae were all fitted by templates showing a prominent second peak, i.e. inconsistent with the intrinsically underluminous supernovae necessary to explain the apparent faintness of high redshift supernovae in a flat $\Lambda = 0$ universe. Note that, the data presented here for SN 2000fr show for the first time a case where the secondary peak is evident in the data even prior to the light curve fit. Table 11 lists the $\Delta\chi^2$ for the fit of the high redshift SNe to the templates of the two underluminous SN 1991bg and SN 1997cn, relative to the best fit. The $\chi^2$ values are significantly larger than the best fit value.

## 8. Summary and conclusions

In this work we have investigated the feasibility of using restframe *I*-band observations for cosmological purposes.

We have developed a five parameter light curve fitting procedure which was applied successfully to 42 nearby type Ia

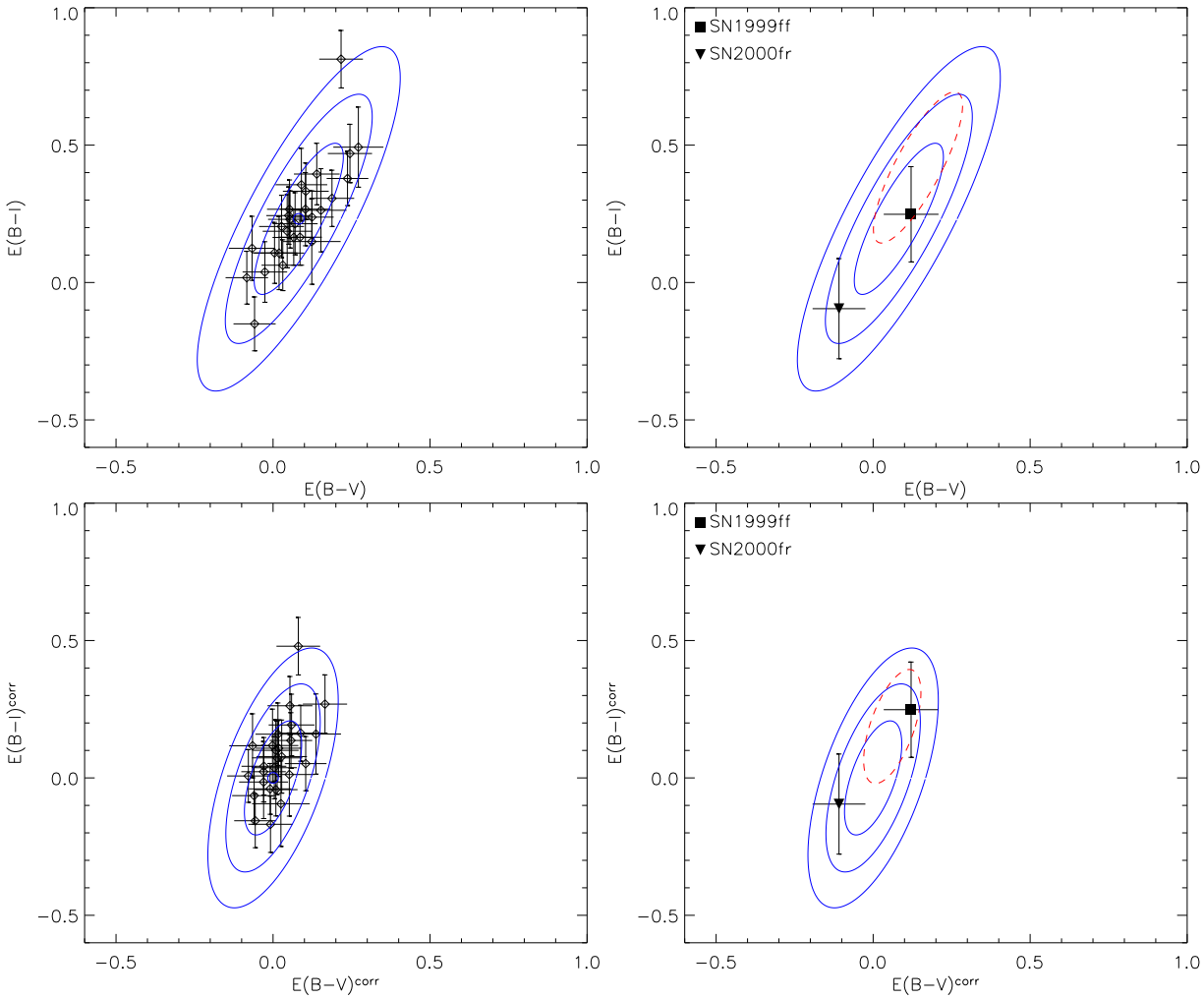


Fig. 13 Left panels: distribution of combined colour measurements for the local sample of supernovae in the $E(B-I)$ vs $E(B-V)$ plane, not corrected (top panel) and corrected (bottom panel) for the host galaxy extinction. The solid contour is 68.3%, 95.5% and 99.7% of probability. Right panels: combined colours of the high redshift supernovae compared to the distribution in the absence of IG dust (solid ellipses), and the one expected in presence of grey dust in the IG medium with $R_V = 9.5$ (dashed ellipse). For simplicity only the 68.3% level has been plotted for the dust distribution.

| | $n$ | SN 1991bg | SN 1997cn |
|---|---|---|---|
| SN 2000fr | 3 | 24.03 | 21.73 |
| SN 1999ff | 2 | 3.28 | 2.72 |

Table 11 $\Delta\chi^2$ for the fit of the high redshift SNe to the templates of the two underluminous SNe relative to the best fits (which are "normal" SN templates). $n$ is the number of data points used in the one-parameter fit (see discussion in section 4.4).

supernovae. The fitted light curve were used to build a set of templates which include a broad variety of shapes. We have found correlations between the fitted parameters, in particular between the time of the secondary peak and the $B$-band stretch, $s_B$. Moreover, a width-luminosity relation was found for the $I$-band peak magnitude both with $s_B$ and with the $I$-band stretch, $s_I$.

We built a restframe $I$-band Hubble diagram using 26 nearby supernovae at redshifts $0.01 \leq z \leq 0.1$, and measured an r.m.s. of 0.24 mag, smaller than the uncorrected dispersion corresponding to restframe $B$-band. The width-luminosity relation was used to reduce the r.m.s. to 0.17 ± 0.03 mag (including measurement errors), corresponding to an intrinsic dispersion of 0.13 mag. Differences between the three data samples are also discussed.

***J*-band measurements of one high redshift supernova plus published data of one more were used to extend the Hubble diagram up to $z \sim 0.5$.** The restframe $I$-band light curves of the $z \sim 0.5$ supernovae were fitted with a template set built on the local SNe sample, as the five parameter fit method could not be used for the poorly sampled high redshift light curve. The peak $I$-band magnitude of the high redshift SNe was compared to three different sets of cosmological parameters. Although the "concordance model" of the universe, $(\Omega_M, \Omega_\Lambda)$=(0.25,0.75), is found in better agreement with the data than the other models, the low statistics of the sample does not allow to reach firm conclusions.

Alternative explanations of the observed dimming of supernova brightness, such as presence of grey dust in the IG

medium or evolution effects in the supernova properties have also been addressed. Both the *I*-band Hubble diagram and multi colour photometry have been used for testing grey dust. Although no firm limits on the presence of grey dust could be set, this study shows that with higher statistics, the restframe *I*-band measurements could provide useful information on cosmological parameters, including tests for systematic effects. A Monte Carlo simulation indicates that a sample of at least 20 well observed SNe would be enough for testing the presence of a homogeneous dust distribution in the IGM, using only the colour diagram technique. A similar technique, but using QSO colours was successfully used by Mörtsell & Goobar (2003), to rule out grey dust being the sole explanation for the apparent faintness of SNe Ia at $z \sim 0.5$.

Possible systematic uncertainties affecting the restframe *I*-band Hubble diagram are discussed. Some sources are identified, for instance the different methods applied for fitting the low and the high redshift samples, selection effects for bright objects during the search campaign, as well as uncertainties in the *K*-correction calculations due to the presence of the Ca IR triplet feature in the near infrared region of the SN spectra. However, these systematic uncertainties differ from the ones that could affect the restframe *B*-band Hubble diagram. Thus, the use of *I*-band measurements of type Ia supernovae, can be complementary to the already well established tools.

*Acknowledgements.* S.N. is grateful to Brian Schmidt for useful discussions on *K*-corrections. Part of this work was supported by a graduate student grant from the Swedish Research Council. AG is a Royal Swedish Academy Research Fellow supported by a grant from the Knut and Alice Wallenberg Foundation. This work was supported in part by the Director, Office of Science, Office of High Energy and Nuclear Physics, of the U.S. Department of Energy under Contract No. DE-AC03-76SF00098. Support for this work was provided by NASA through grant HST-GO-08346.01-A from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5-26555.

## References

Aldering, G.,2000, coex.conf, 75A
Aguirre, A. 1999, ApJ, 512, L19
Aguirre, A. 1999, ApJ, 525, 583
Balbi, A., Ade, P., Bock, J. et al., 2000, ApJ, 545, L1.
Bessell, M.S., 1990, PASP, 102, 1181B
Bessell, M.S., Castelli, F. & Plez, B. 1998, A&A, 333, 231
Borgani, S., Rosati, P., Tozzi, P. et al., 2001, ApJ, 561, 13
Cardelli, J. A, Clayton, G.C. & Mathis, J.S. 1989, ApJ, 345, 245
Carpenter, J.M., 2001, AJ, 121, 2851C
Cohen, M., Wheaton, WM. A. & Megeath, AJ, 2003, 126,1090
Contardo, G., Leibundgut, B. & Vacca, W. D. 2000, A&A , 359, 876C
Cowan, G., 1998, Statistical data analysis, Oxford University Press
Csaki, C., Kaloper, N. & Terning, J., 2002, Phys.Rev.Lett. 88, 161302
De Bernardis, P., Ade, P.A.R., Bock, J.J. et al., 2000, Nature, 404,955
Deffayet, C., Harari, D., Uzan, J.P. & Zaldarriaga, M., 2002, Phys.Rev.D66, 043517
Drell, P., Loredo, T.& Wasserman, I., 2000, ApJ, 530, 593.
Efstathiou, G., Moody, S., Peacock, J.A. et al., 2002, MNRAS, 330, L29.
Filippenko, A.V., Richmond, M.W., Matheson, T., 1992a, ApJ, 384L, 15F
Filippenko, A.V., Richmond, M.W., Branch, D. et al., 1992b, AJ,104,1543
Garavini, G. et al., 2004a (in prep)
Garavini, G., Folatelli, G., Goobar, A. et al., 2004b, AJ, 128, 387G
Garnavich, P.M., Kirshner, R.P., Challis, P., 1998, ApJ, 493L, 53G
Garnavich, P.M. et al., 1999,IAUC, 7097
Goldhaber, G., Groom, D. E.,Kim, A. et al., 2001, ApJ, 558, 359G
Goobar, A., Bergström, L. & Mörtsell, E., 2002, A&A, 384, 1.
Goobar, A., Mörtsell, E., Amanullah, R., Goliath, M., Bergström, L. & T. Dahlen, 2002, A&A, 392, 757
Hamuy, M., Maza, J., Wischnjewsky, M. et al., 1993, IAUC, 5723R
Hamuy, M., Phillips, M.M., Suntzeff, N.B. et al. 1996, AJ, 112, 2408
Hamuy, M., Phillips, M.M., Suntzeff, N.B. et al. 1996, AJ, 112, 2398H
Henry J. P. 2001, ApJ 534, 565
Howell, D.A, 2001, ApJ,554, 193
Jaffe, A.H., Ade, P.A., Balbi, A. et al., 2001, Phys. Rev. Lett. 86, 3475
Jha, S., 2002, PhD thesis, Harvard University.
Kim, A., Goobar, A. & Perlmutter, S., 1996, PASP, 108, 190
Knop, R. et al., 2003, submitted
Krisciunas, K., Phillips, M.M., Stubbs, C. et al., 2001, AJ, 122, 1616K
Krisciunas, K., Phillips, M. M. & Suntzeff, N. B., astro-ph/0312626
Leibundgut, B., Kirshner, R. P., Phillips, M.M. et al., 1993,AJ, 105, 301.
Li, W., Filippenko, A.V. & Treffers, R.R., 2001, ApJ, 546,734
Lidman, C. et al., 2004 (in prep)
Maza, J., Hamuy, M., Phillips, M., Suntzeff, N. & Aviles, R., 1994, ApJ,424,L107
Mörtsell, E., Bergström, L and Goobar, A., 2002, Phys. Rev. D66, 047702.
Mörtsell, E. & Goobar, A., 2003, JCAP, 09, 009M
Nobili, S., Goobar, A., Knop, R., & Nugent, P., 2003, A&A 404, 901-912
Nugent, P. & Aldering, G., 2000, sgrb.conf, 47N
Nugent P., Kim A.& Perlmutter S., 2002, PASP 114, 803
Östman, L. & Mörtsell, E., 2004, astro-ph/040501
Perlmutter, S., Gabi, S., Goldhaber, G. et al., 1997,ApJ, 483, 565
Perlmutter, S., Aldering, G., della Valle, M. et al., 1998, Nature, 391, 51P
Perlmutter, S., Aldering, G., Goldhaber, G. et al., 1999, ApJ, 517, 565

Persson, S.E., Murphy, D.C., Krzeminski, W., Roth, M. & Rieke, M.J., 1998 A.J. 116, 2475.
Phillips, M.M., Wells, L.A., Suntzeff, N.B. et al. 1992, AJ, 103,1632P
Phillips, M.M., Lira, P., Suntzeff, N.B., Schommer, R.A., Hamuy, M.& Maza, J., 1999, AJ, 118:1766-1776
Phillips, M. M.; Krisciunas, K.; Suntzeff, N. B. et al., 2003, fthp.conf, 193P
Richmond, M.W., Treffers, R.R., Filippenko, A.V. et al., 1995, AJ 109,2121
Riess, A., Filippenko, A.V., Challis, P. et al., 1998, AJ, 116, 1009.
Riess, A.G., Kirshner, R.P.,Schmidt, B.P. et al., 1999, ApJ. 117,707-724
Riess, A.G., Filippenko, A.V., Liu, M.C. et al., 2000, ApJ. 536,62-67
Riess, A.G., Strolger, L-G.; Tonry, J. et al., 2004, ApJ, 607, 665R
Rowan-Robinson, M.,2002,MNRAS, 332, 352.
Ruiz-Lapuente, P., Cappellaro, E., Turatto, M. et al. 1992, ApJ, 387L, 33R
Schahmaneche, K. et al., SCP Collaboration ,IAUC 7763 (2001) 1
Schmidt, B.P., Suntzeff, N.B., Phillips, M.M. et al., 1998, ApJ, 507, 46S
Schlegel, D.J., Finkbeiner, D.P. & Davis, M. 1998, ApJ, 500, 525S
Schuecker, P. Böhringer, H., Collins, C. A. & Guzzo, L., 2003, A&A, 398, 867
Sievers, J.L., Bond, J.R., Cartwright, J.K. et al., 2003, ApJ.,591, 599
Spergel, D.N., Verde, L., Peiris, H.V. et al., 2003, ApJS, 148, 175S
Strolger, L.-G., Smith, R. C., Suntzeff, N. B. et al., 2002, AJ, 124, 2905
,Sullivan, M., Ellis, R. S., Aldering, G. et al., 2003, MNRAS, 340, 1057S
Tonry, J.L.,Schmidt B.P., Barris, B. et al., 2003, ApJ, 594, 1T
Wang, L., Goldhaber, G., Aldering, G., Perlmutter, S.2003, ApJ, 590, 944
Wang, L., Conley, A., Goldhaber, G., et al., 2005 (in preparation)
Wells, L.A., Phillips, M.M., Suntzeff, B. et al., 1994 AJ 108, 2233